

Scientists find most Earth-like planet yet

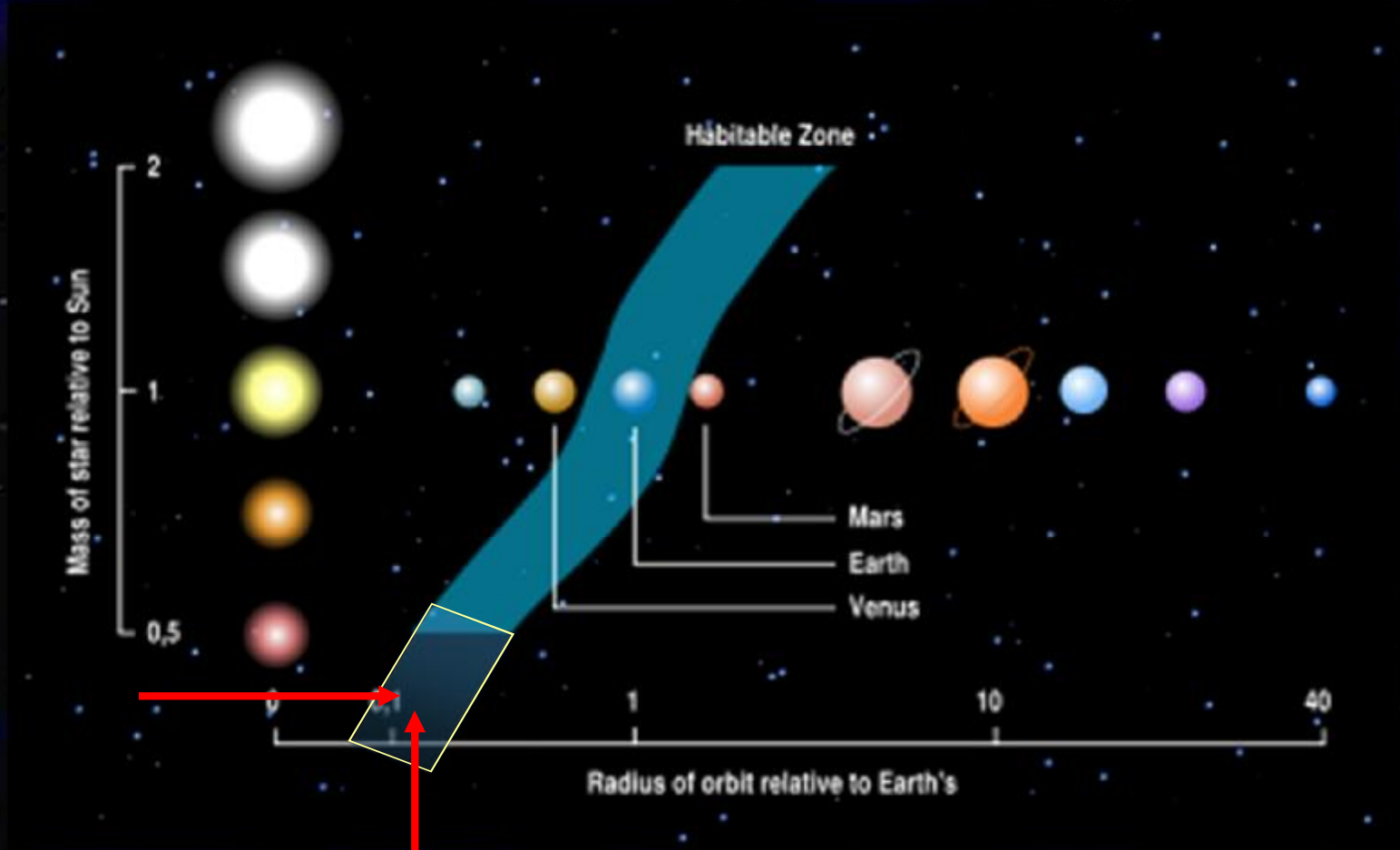
Models predict planet should be either rocky or covered with oceans

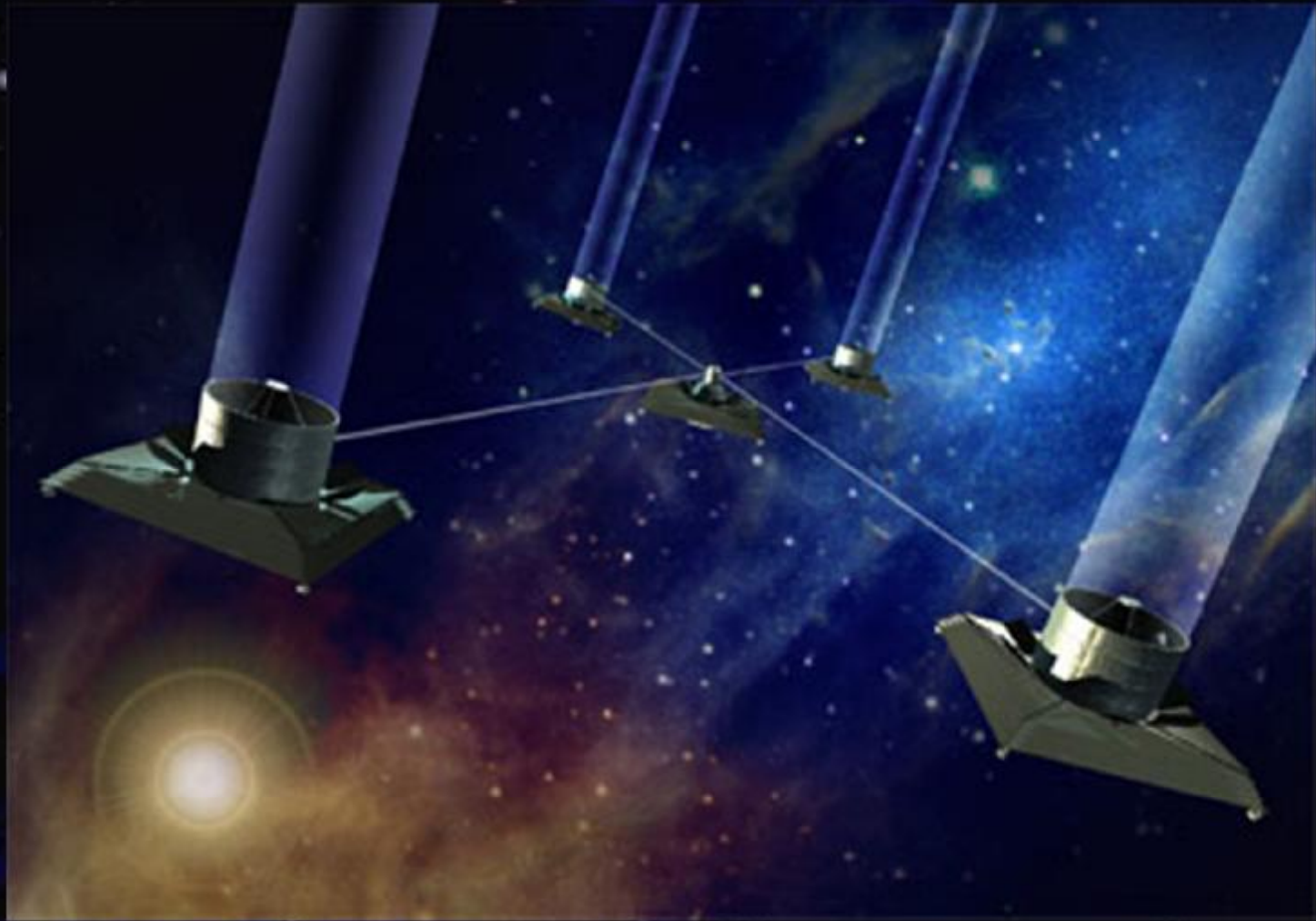


Gliese 581

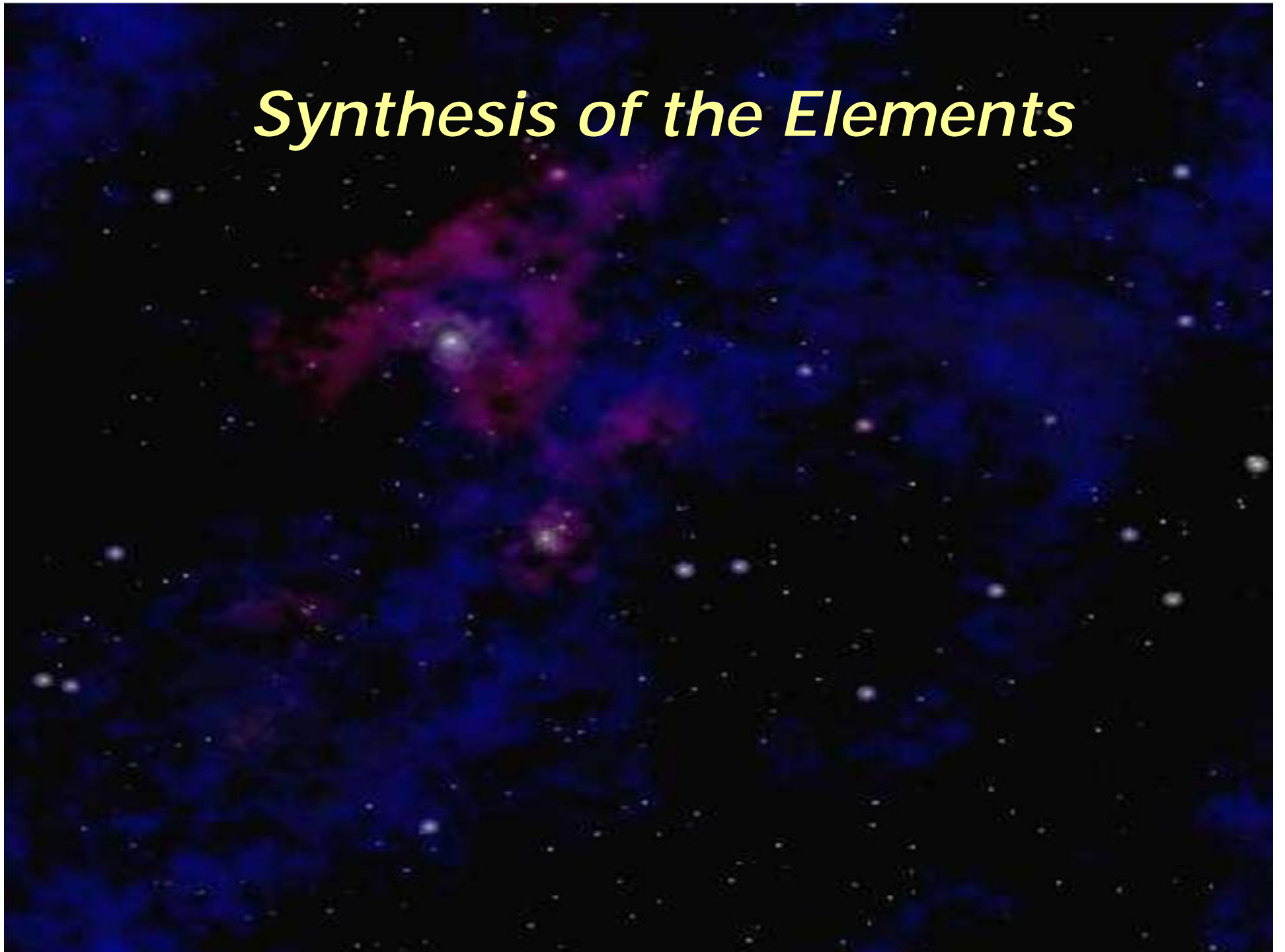
50% larger than earth, 5 times the mass
150 lb person would weight 333 lbs
Temperature: 32-104°F

6,000,000 miles from M type star
13 days to complete orbit

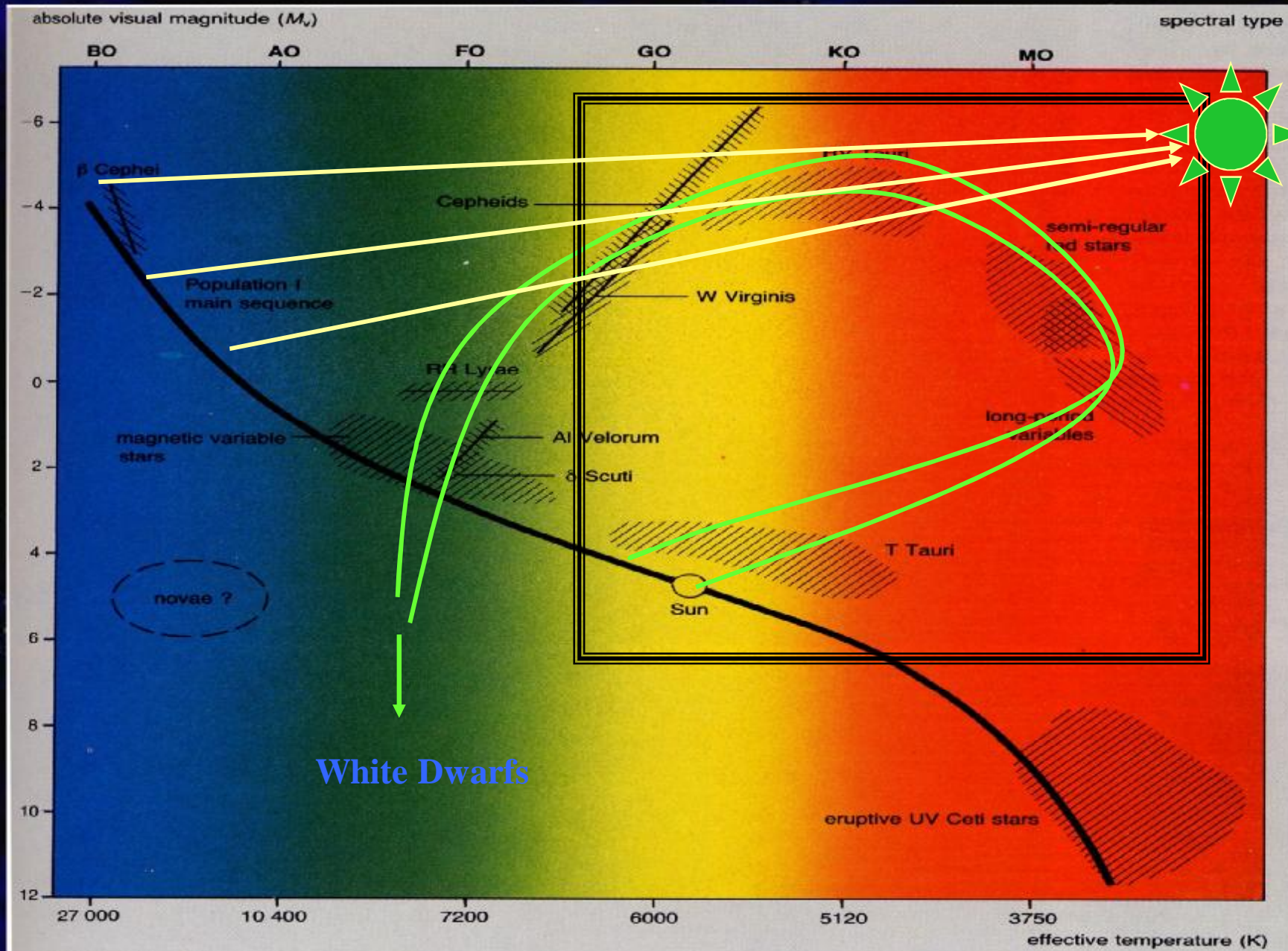




Synthesis of the Elements



1 <u>H</u> 1.008	2 IIA 2A <u>Be</u> 9.012											13 IIIA 3A	14 IVA 4A	15 VA 5A	16 VIA 6A	17 VIIA 7A	18 VIII 8 9 10 11 12 IB 1B 2B	19 IIA 2A <u>Ca</u> 40.08	20 IIIB 3B <u>Sc</u> 44.96	21 IVB 4B <u>Ti</u> 47.88	22 VB 5B <u>V</u> 50.94	23 VIB 6B <u>Cr</u> 52.00	24 VIIB 7B <u>Mn</u> 54.94	25 VIII 8 <u>Fe</u> 55.85	26 VIII 8 <u>Co</u> 58.47	27 VIII 8 <u>Ni</u> 58.69	28 VIII 8 <u>Cu</u> 63.55	29 VIII 8 <u>Zn</u> 65.39	30 VIII 8 <u>Ga</u> 69.72	31 VIII 8 <u>Ge</u> 72.59	32 VIII 8 <u>As</u> 74.92	33 VIII 8 <u>Se</u> 78.96	34 VIII 8 <u>Br</u> 79.90	35 VIII 8 <u>Kr</u> 83.80	
3 <u>Li</u> 6.941	4 <u>Be</u> 9.012	3 IIIB 3B <u>Na</u> 22.99	4 IVB 4B <u>Mg</u> 24.31	5 VB 5B	6 VIB 6B	7 VIIB 7B	8 VIII 8	9 VIII 8	10 VIII 8	11 IB 1B	12 IIB 2B	13 IIIA 3A <u>Al</u> 26.98	14 IVA 4A <u>Si</u> 28.09	15 VA 5A <u>P</u> 30.97	16 VIA 6A <u>S</u> 32.07	17 VIIA 7A <u>Cl</u> 35.45	18 VIII 8 <u>Ar</u> 39.95	19 IIA 2A <u>K</u> 39.10	20 IIIB 3B <u>Ca</u> 40.08	21 IVB 4B <u>Sc</u> 44.96	22 VB 5B <u>Ti</u> 47.88	23 VIB 6B <u>V</u> 50.94	24 VIIB 7B <u>Cr</u> 52.00	25 VIII 8 <u>Mn</u> 54.94	26 VIII 8 <u>Fe</u> 55.85	27 VIII 8 <u>Co</u> 58.47	28 VIII 8 <u>Ni</u> 58.69	29 VIII 8 <u>Cu</u> 63.55	30 VIII 8 <u>Zn</u> 65.39	31 VIII 8 <u>Ga</u> 69.72	32 VIII 8 <u>Ge</u> 72.59	33 VIII 8 <u>As</u> 74.92	34 VIII 8 <u>Se</u> 78.96	35 VIII 8 <u>Br</u> 79.90	36 VIII 8 <u>Kr</u> 83.80
37 <u>Rb</u> 85.47	38 <u>Sr</u> 87.62	39 <u>Y</u> 88.91	40 <u>Zr</u> 91.22	41 <u>Nb</u> 92.91	42 <u>Mo</u> 95.94	43 <u>Tc</u> (98)	44 <u>Ru</u> 101.1	45 <u>Rh</u> 102.9	46 <u>Pd</u> 106.4	47 <u>Ag</u> 107.9	48 <u>Cd</u> 112.4	49 <u>In</u> 114.8	50 <u>Sn</u> 118.7	51 <u>Sb</u> 121.8	52 <u>Te</u> 127.6	53 <u>I</u> 126.9	54 <u>Xe</u> 131.3	37 IIA 2A <u>Rb</u> 85.47	38 IIIB 3B <u>Sr</u> 87.62	39 IVB 4B <u>Y</u> 88.91	40 VB 5B <u>Zr</u> 91.22	41 VIB 6B <u>Nb</u> 92.91	42 VIIB 7B <u>Mo</u> 95.94	43 VIII 8 <u>Tc</u> (98)	44 VIII 8 <u>Ru</u> 101.1	45 VIII 8 <u>Rh</u> 102.9	46 VIII 8 <u>Pd</u> 106.4	47 VIII 8 <u>Ag</u> 107.9	48 VIII 8 <u>Cd</u> 112.4	49 VIII 8 <u>In</u> 114.8	50 VIII 8 <u>Sn</u> 118.7	51 VIII 8 <u>Sb</u> 121.8	52 VIII 8 <u>Te</u> 127.6	53 VIII 8 <u>I</u> 126.9	54 VIII 8 <u>Xe</u> 131.3
55 <u>Cs</u> 132.9	56 <u>Ba</u> 137.3	57 <u>La*</u> 138.9	72 <u>Hf</u> 178.5	73 <u>Ta</u> 180.9	74 <u>W</u> 183.9	75 <u>Re</u> 186.2	76 <u>Os</u> 190.2	77 <u>Ir</u> 190.2	78 <u>Pt</u> 195.1	79 <u>Au</u> 197.0	80 <u>Hg</u> 200.5	81 <u>Tl</u> 204.4	82 <u>Pb</u> 207.2	83 <u>Bi</u> 209.0	84 <u>Po</u> (210)	85 <u>At</u> (210)	86 <u>Rn</u> (222)	55 IIA 2A <u>Cs</u> 132.9	56 IIIB 3B <u>Ba</u> 137.3	57 IVB 4B <u>La*</u> 138.9	72 VB 5B <u>Hf</u> 178.5	73 VIB 6B <u>Ta</u> 180.9	74 VIIB 7B <u>W</u> 183.9	75 VIII 8 <u>Re</u> 186.2	76 VIII 8 <u>Os</u> 190.2	77 VIII 8 <u>Ir</u> 190.2	78 VIII 8 <u>Pt</u> 195.1	79 VIII 8 <u>Au</u> 197.0	80 VIII 8 <u>Hg</u> 200.5	81 VIII 8 <u>Tl</u> 204.4	82 VIII 8 <u>Pb</u> 207.2	83 VIII 8 <u>Bi</u> 209.0	84 VIII 8 <u>Po</u> (210)	85 VIII 8 <u>At</u> (210)	86 VIII 8 <u>Rn</u> (222)

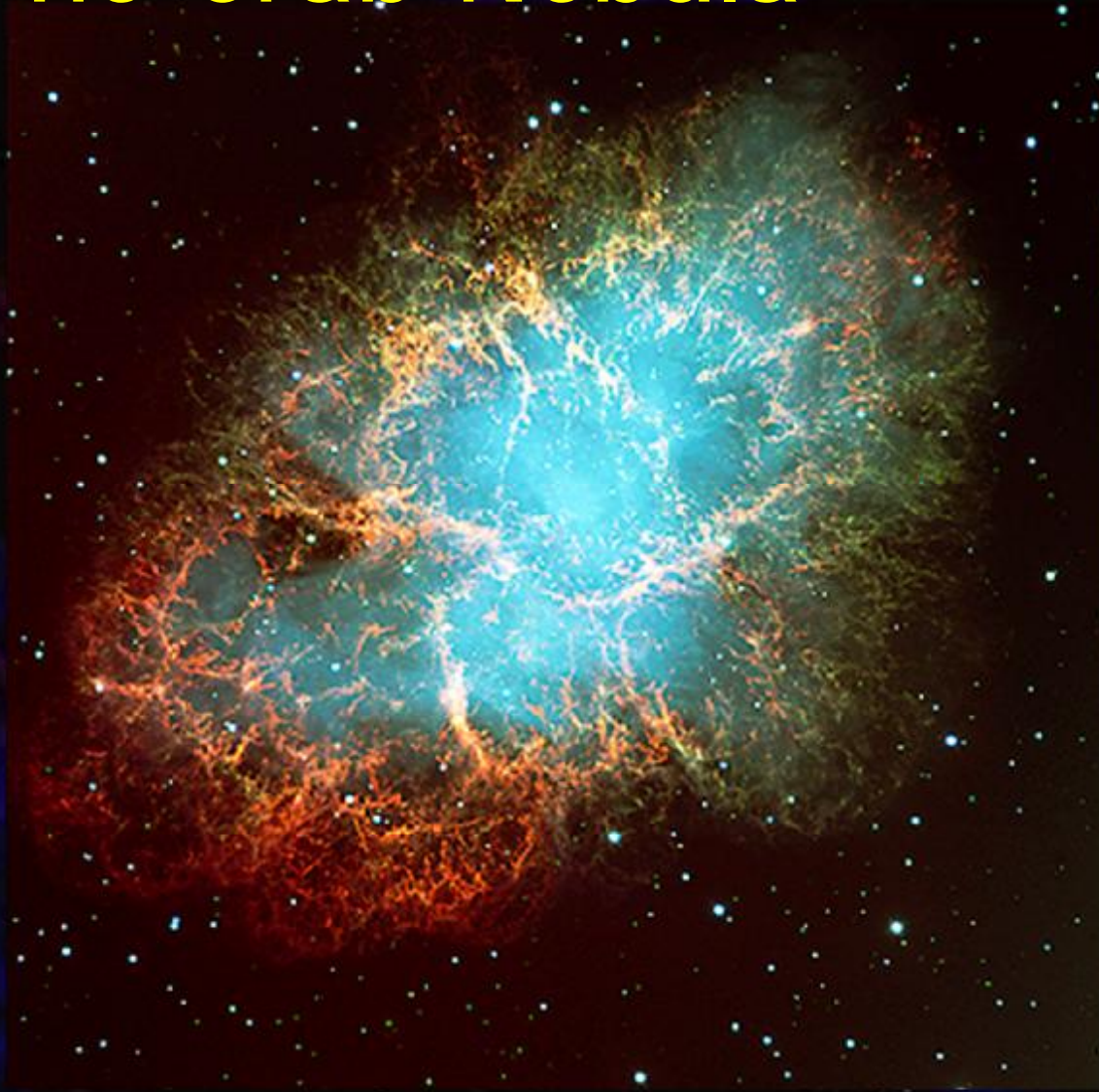


White Dwarfs

July, 1054 A.D.

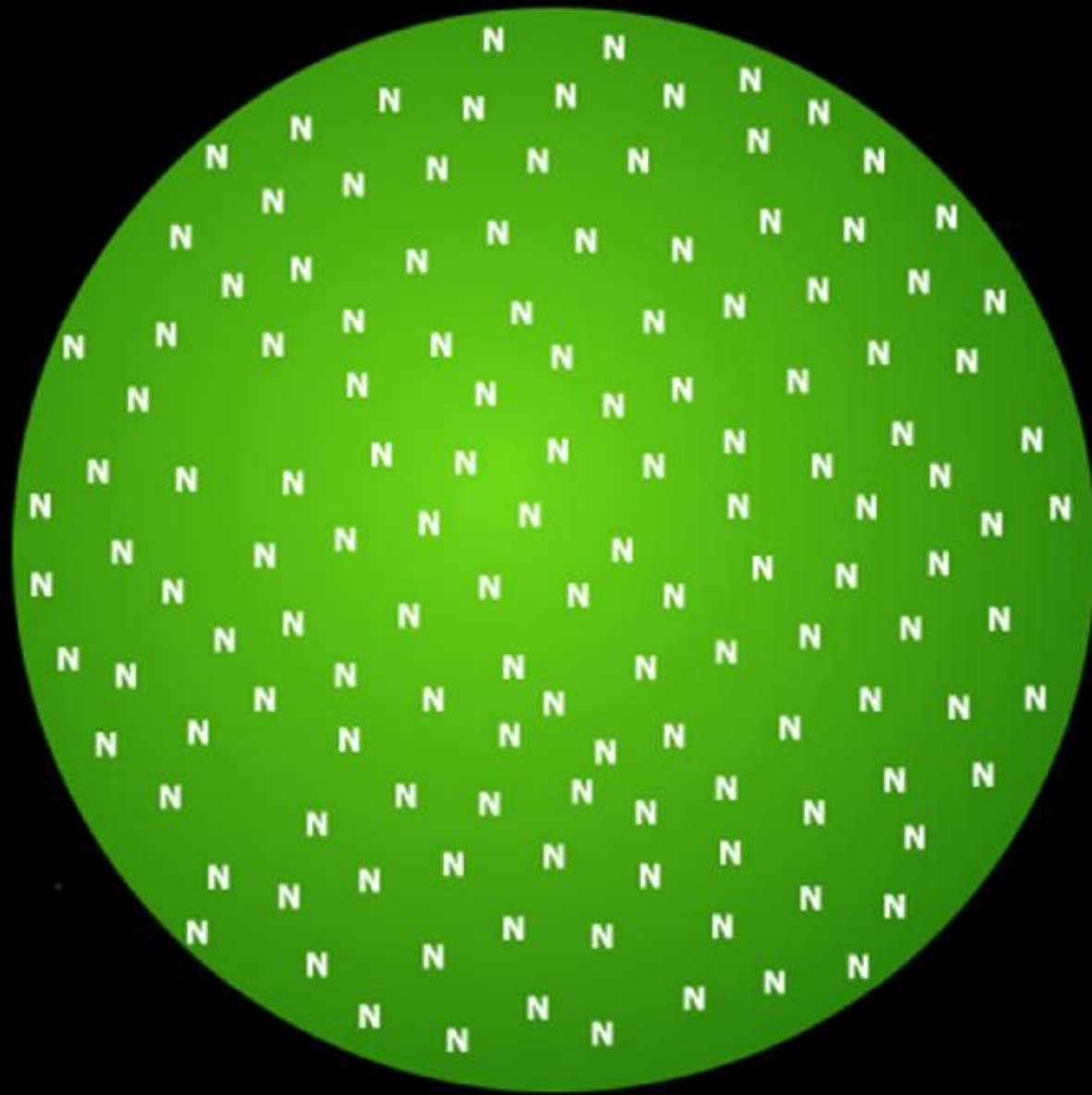


M1 - The Crab Nebula



Crab Nebula

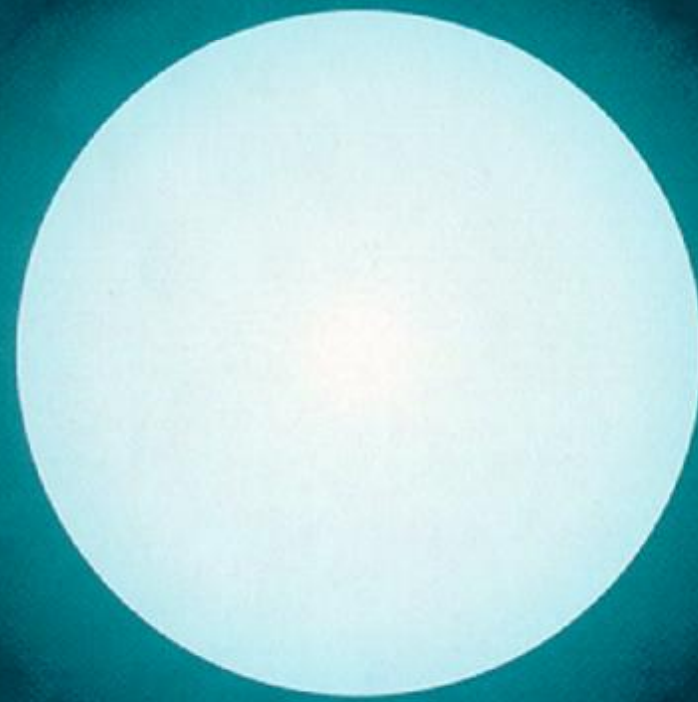




NEUTRON STAR



Earth

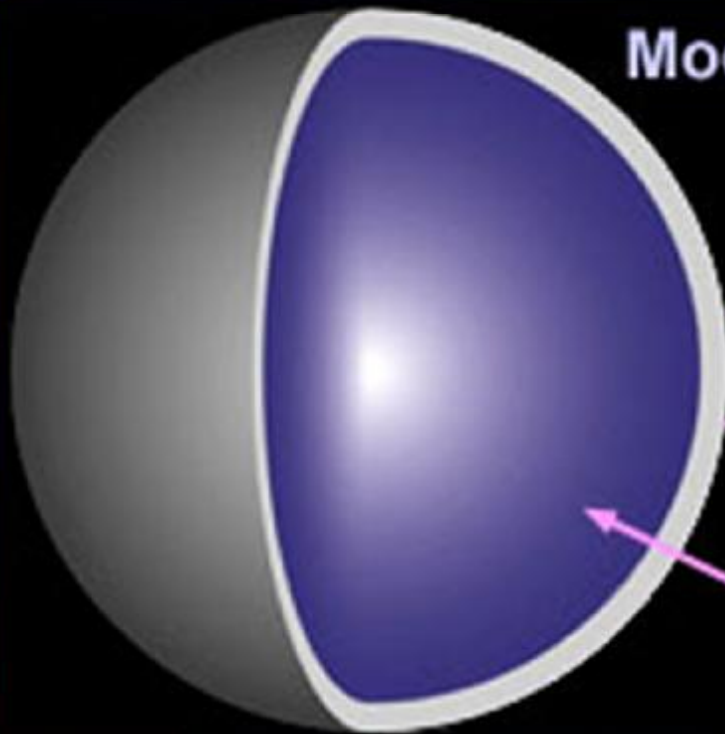


Solar-mass white dwarf



neutron star

Model of a Neutron Star



Mass

~1.5 times the Sun

Solid crust

~1 mile thick

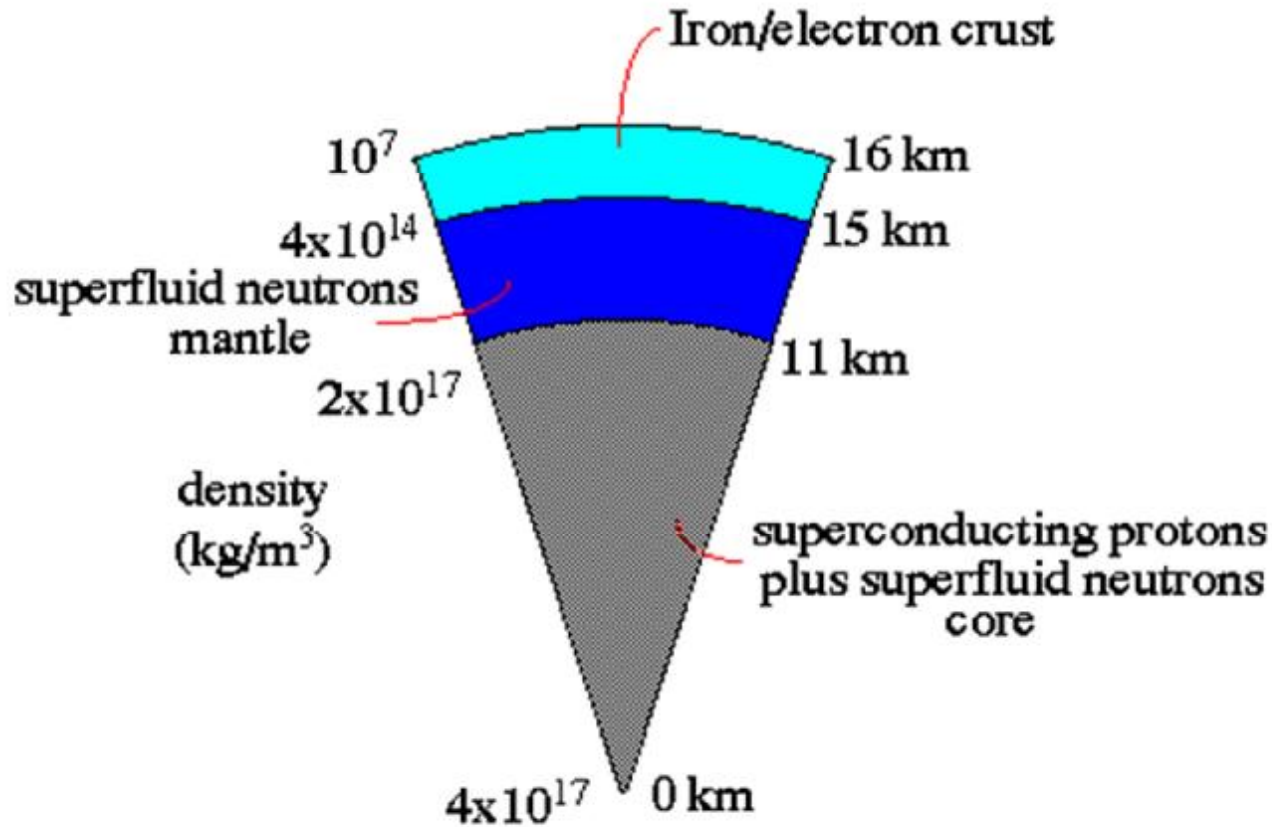
Diameter

~12 miles

Heavy liquid interior

Mostly neutrons,
with other particles

Neutron Star Interior

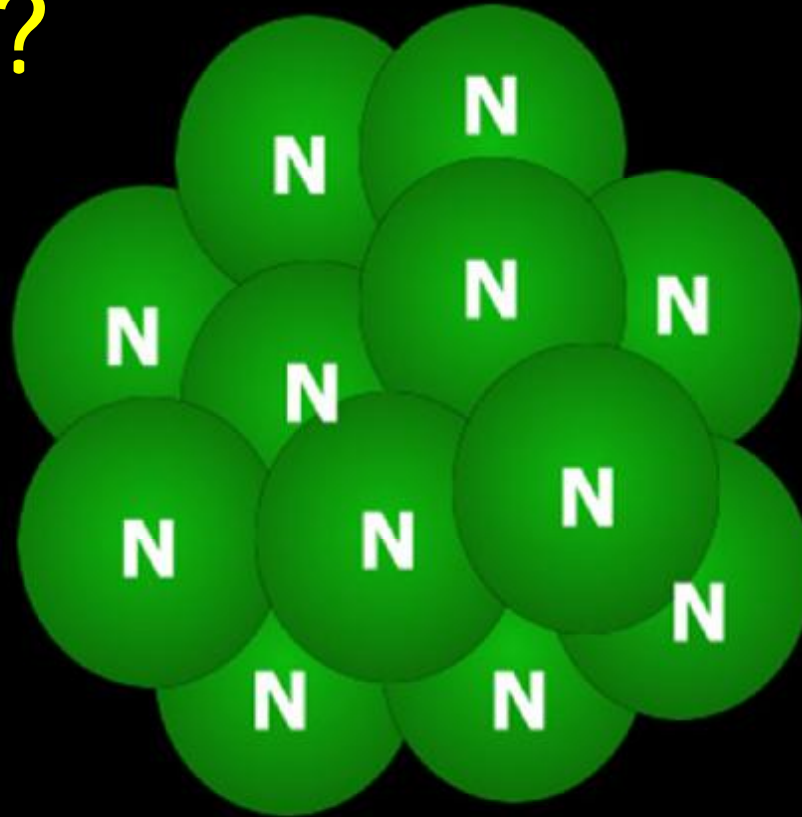


220,000,000 tons
per cubic inch

Limit ~ 3 solar M

Spin up to 38000 rpm
Mag field: 10^{14} gauss

What keeps the neutron star from collapsing?



NEUTRON DEGENERACY PRESSURE

Neutrons have a limit to how tightly they can be packed together

Chandrasekhar Limit for neutron stars

$$M < 3.0 M_{\odot}$$

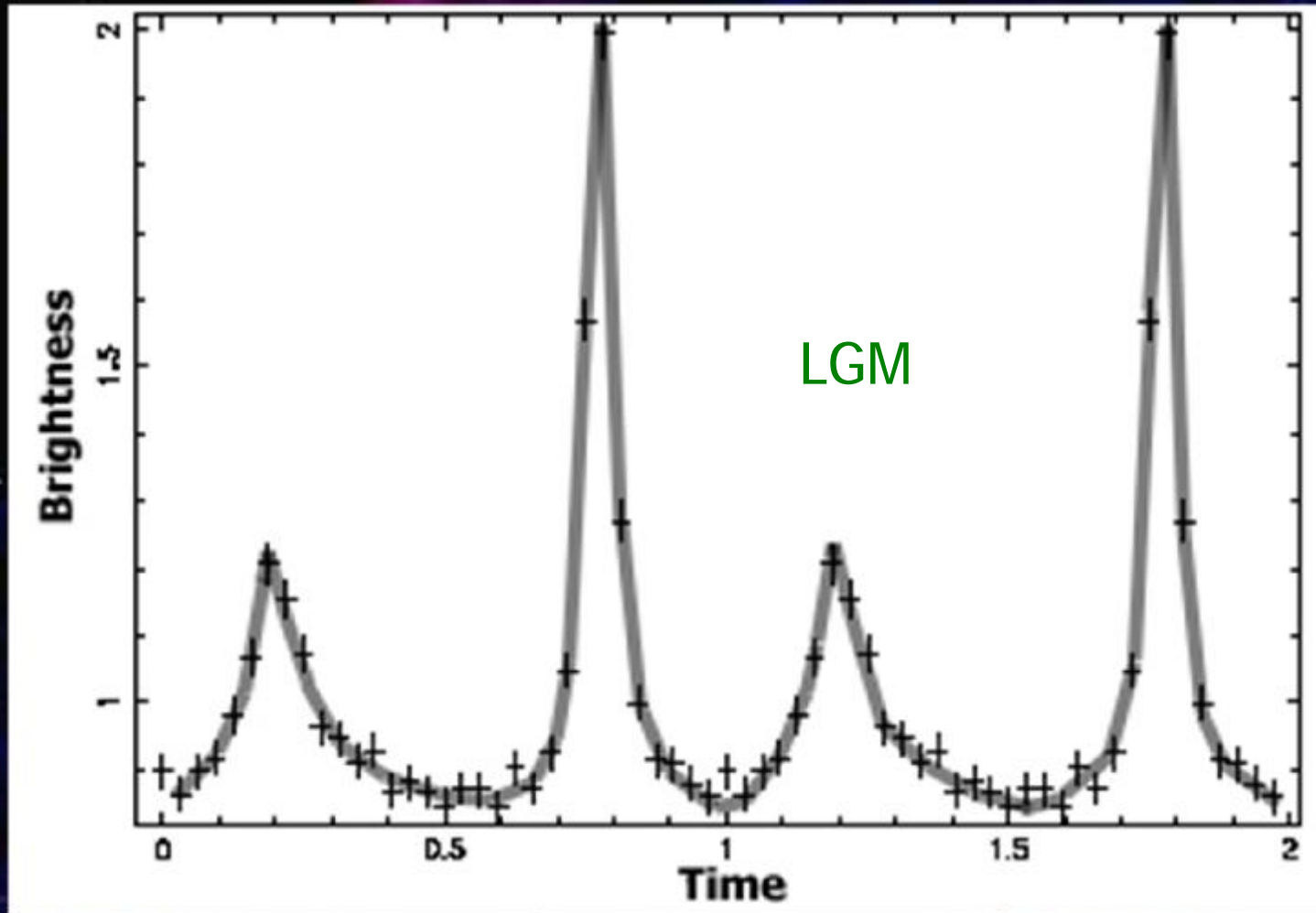
How do we know that neutron stars actually exist?

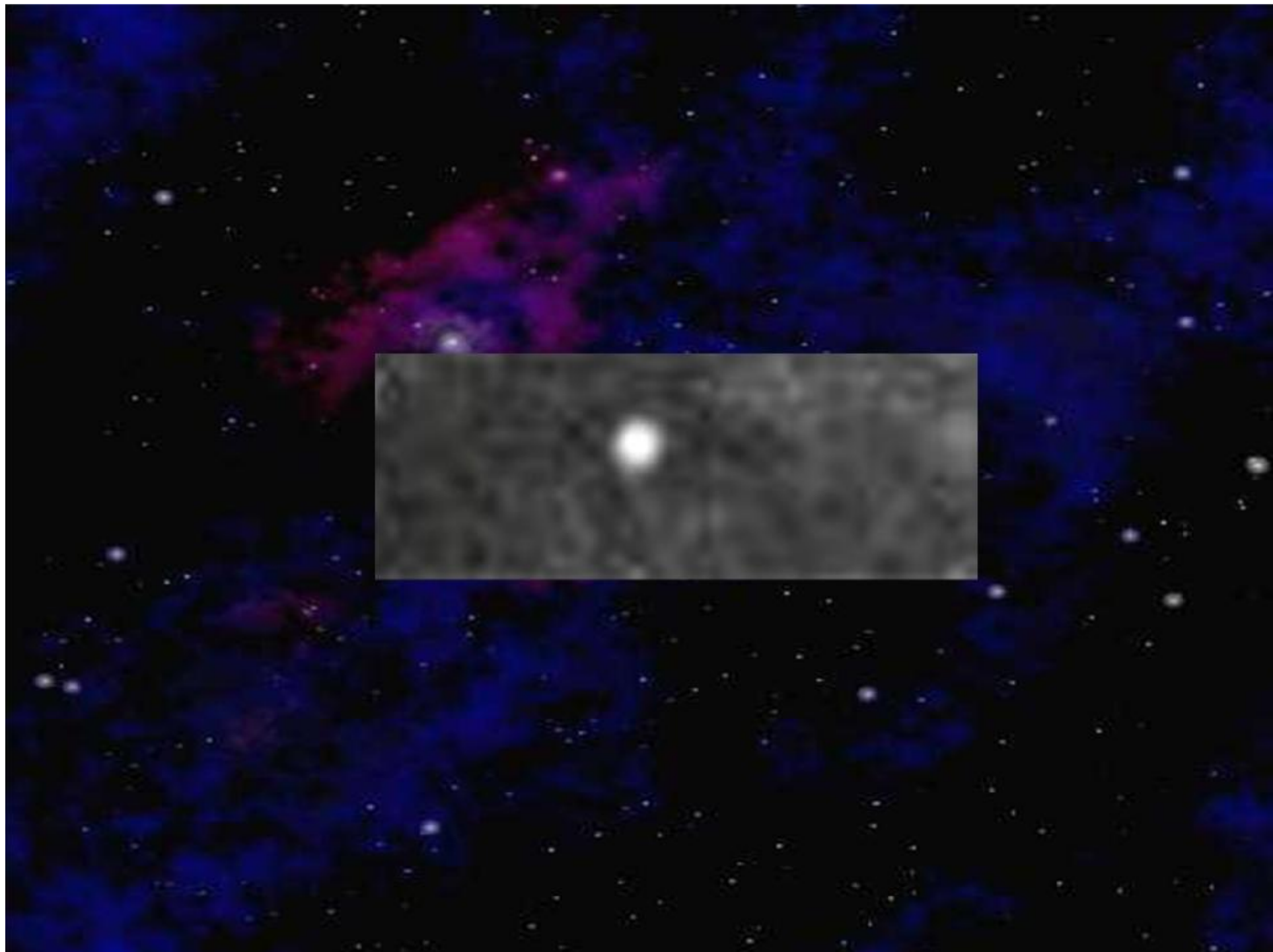
- n First theorized in the 1930's
- n First discovered in 1967

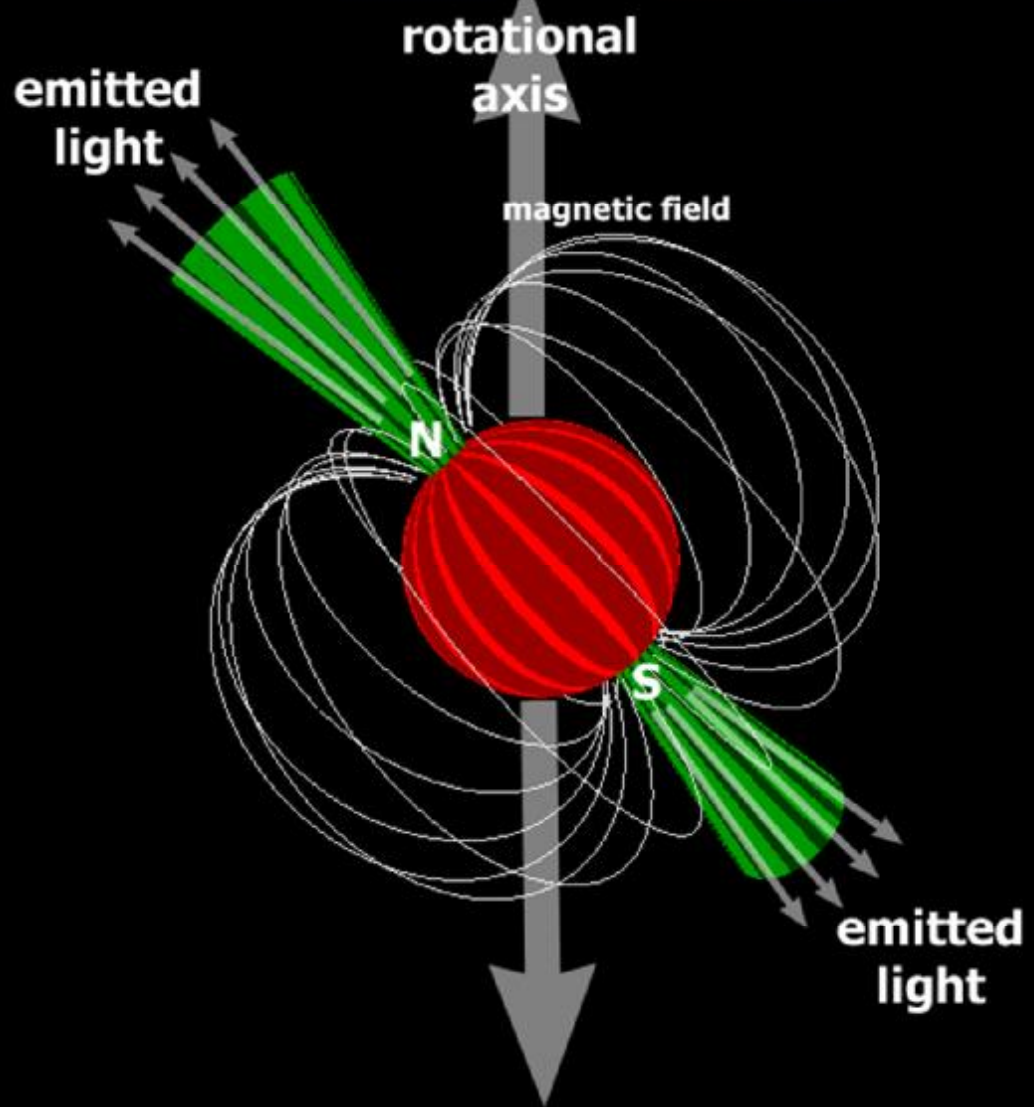




“pulsed” energy every 1.34 seconds

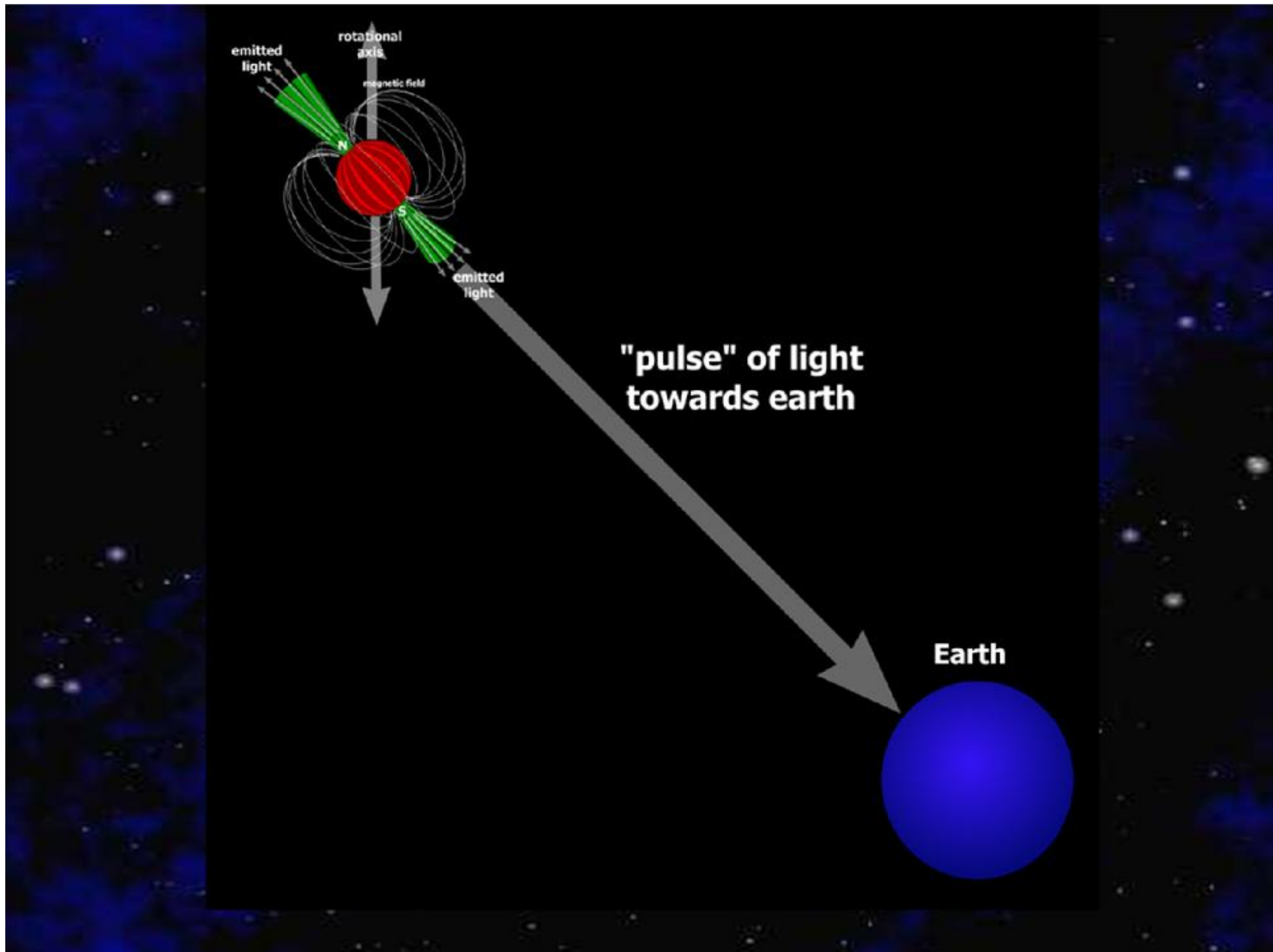






Conservation of Angular Momentum





MPIfR-Bonn Pulsar Group





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Rapidly rotating neutron star

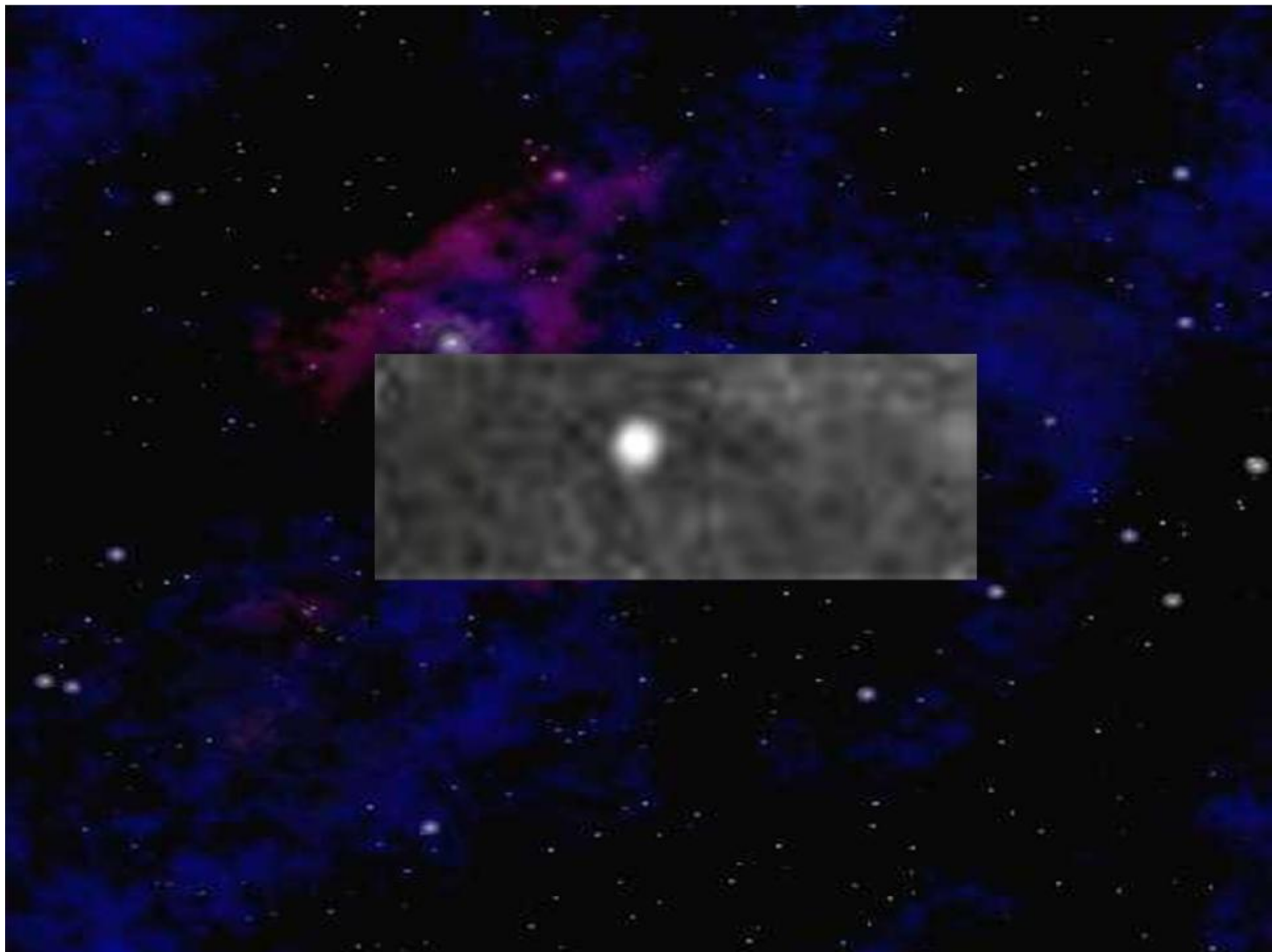
-or-

PULSAR

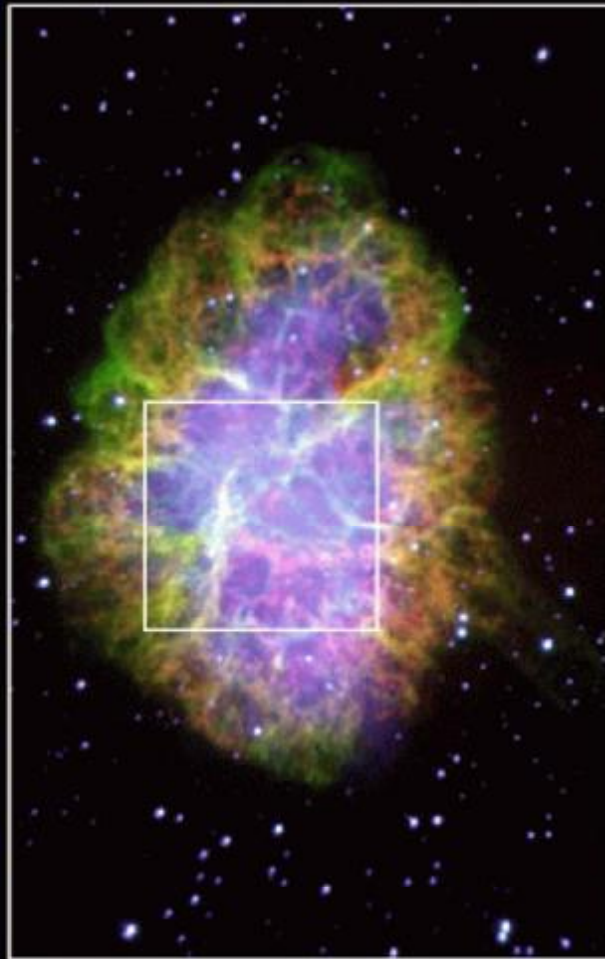
The Crab Nebula







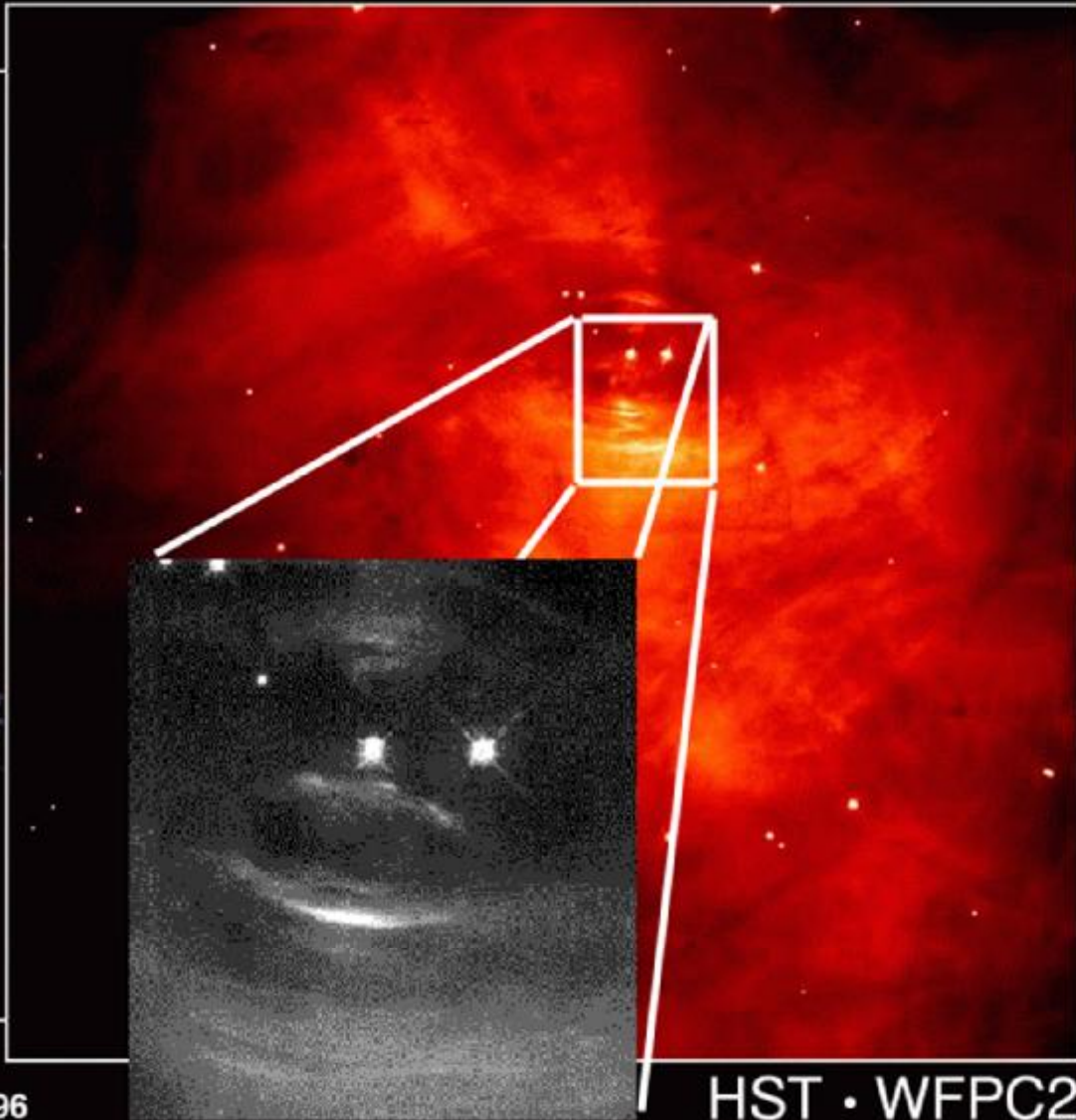
Crab Nebula



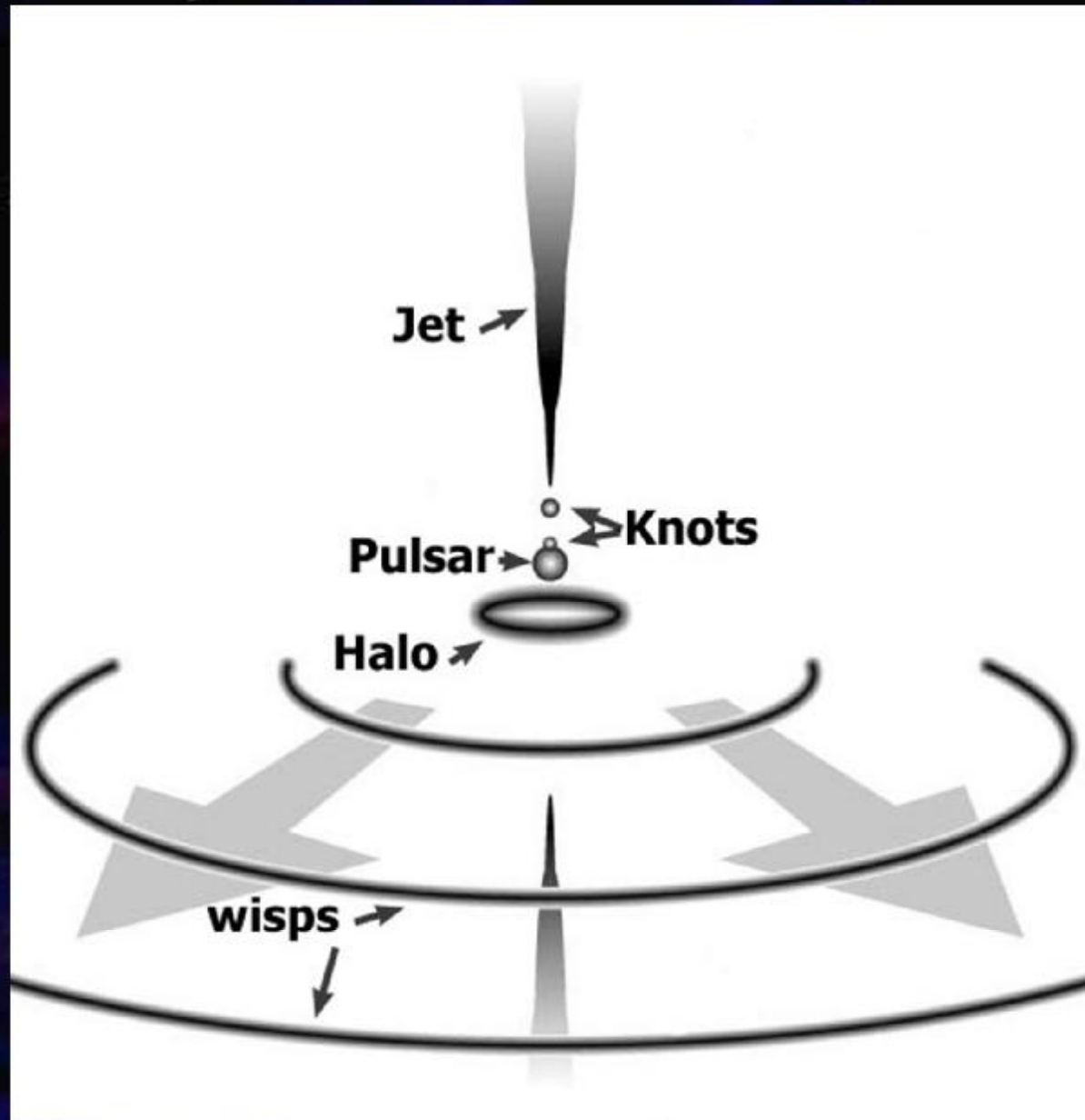
Palomar

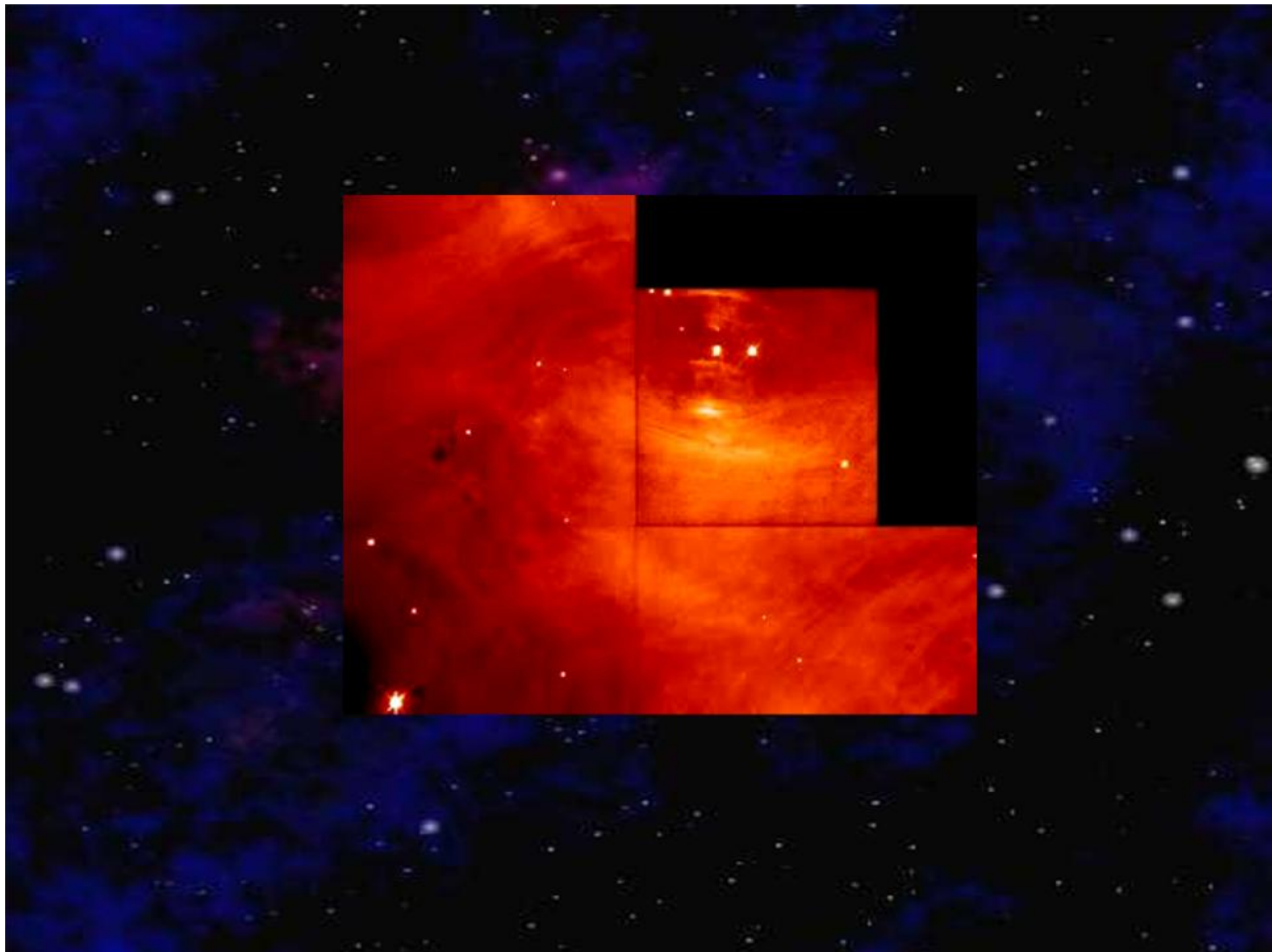
PRC96-22a · ST ScI OPO · May 30, 1996

J. Hester and P. Scowen (AZ State Univ.) and NASA

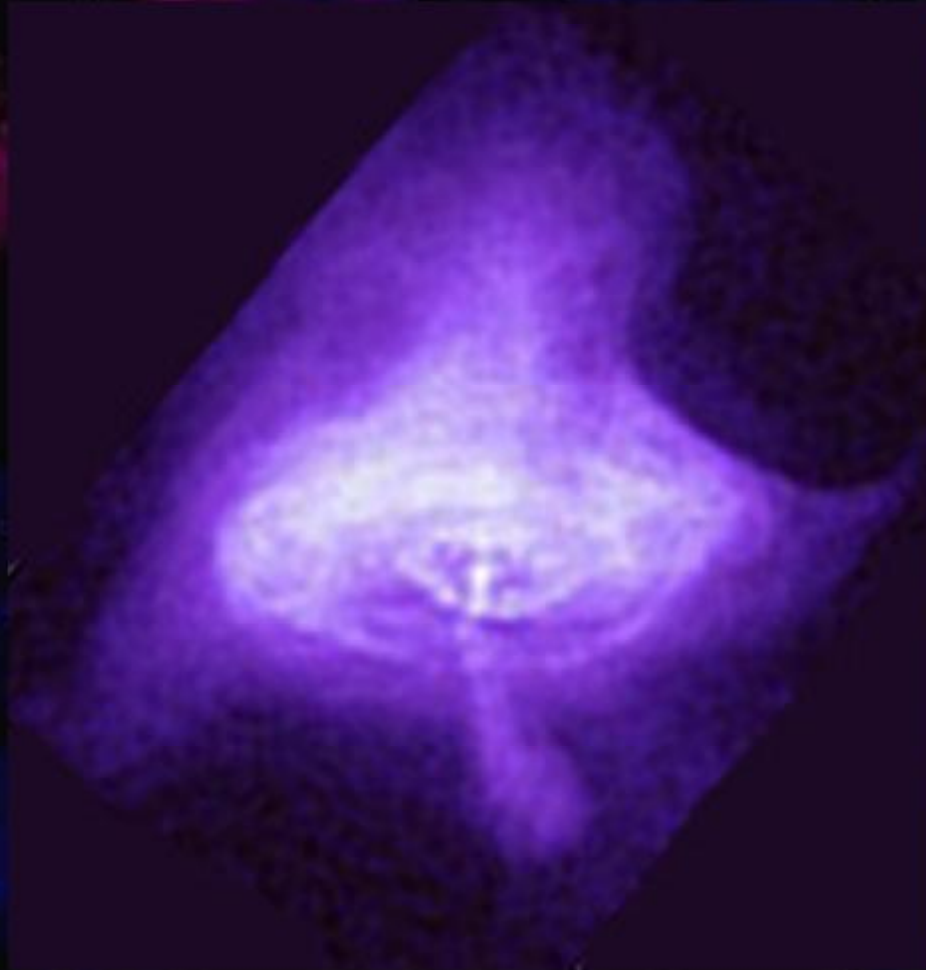


HST · WFPC2





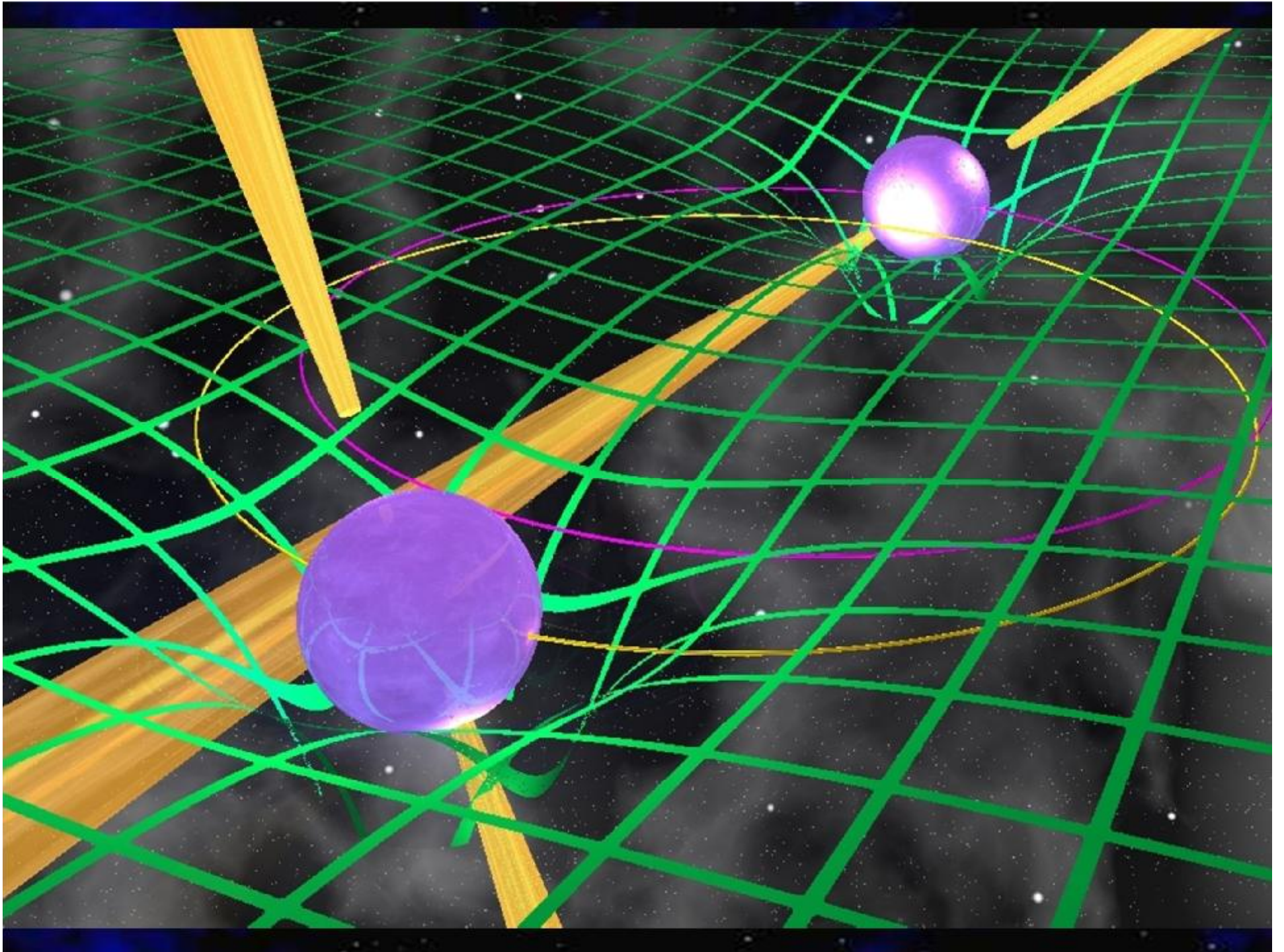
The Crab Pulsar – Xray image

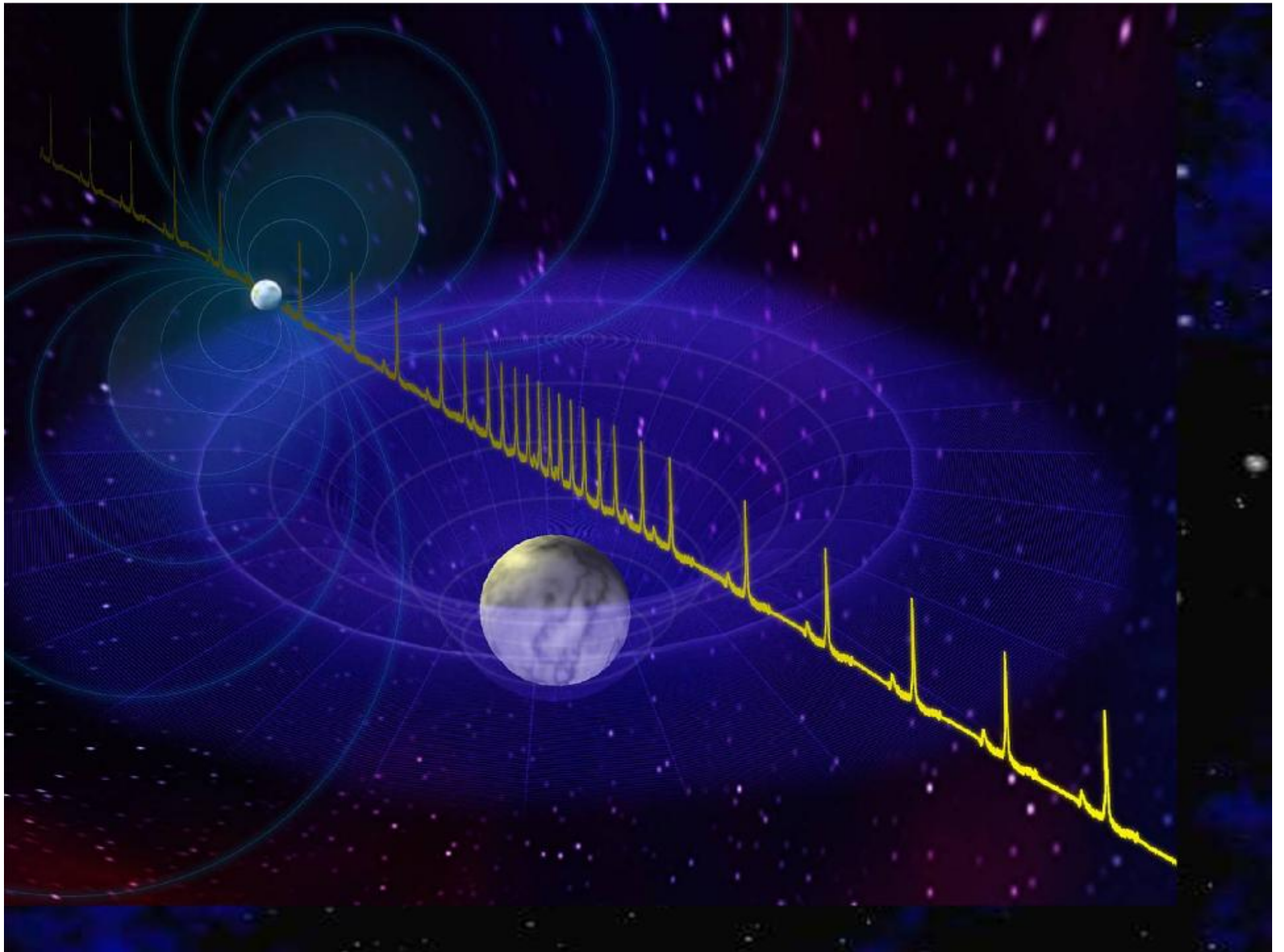


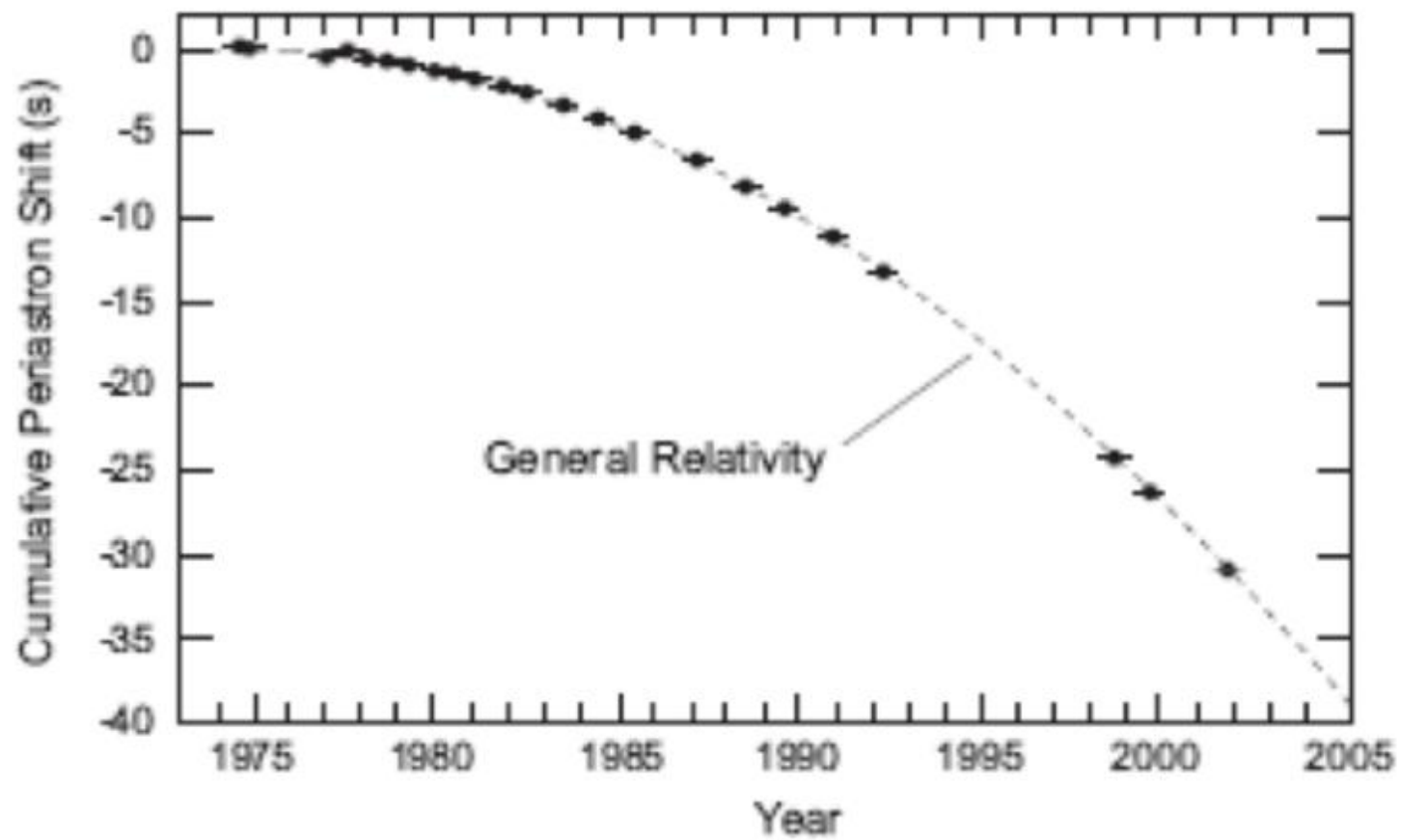
The Crab Pulsar

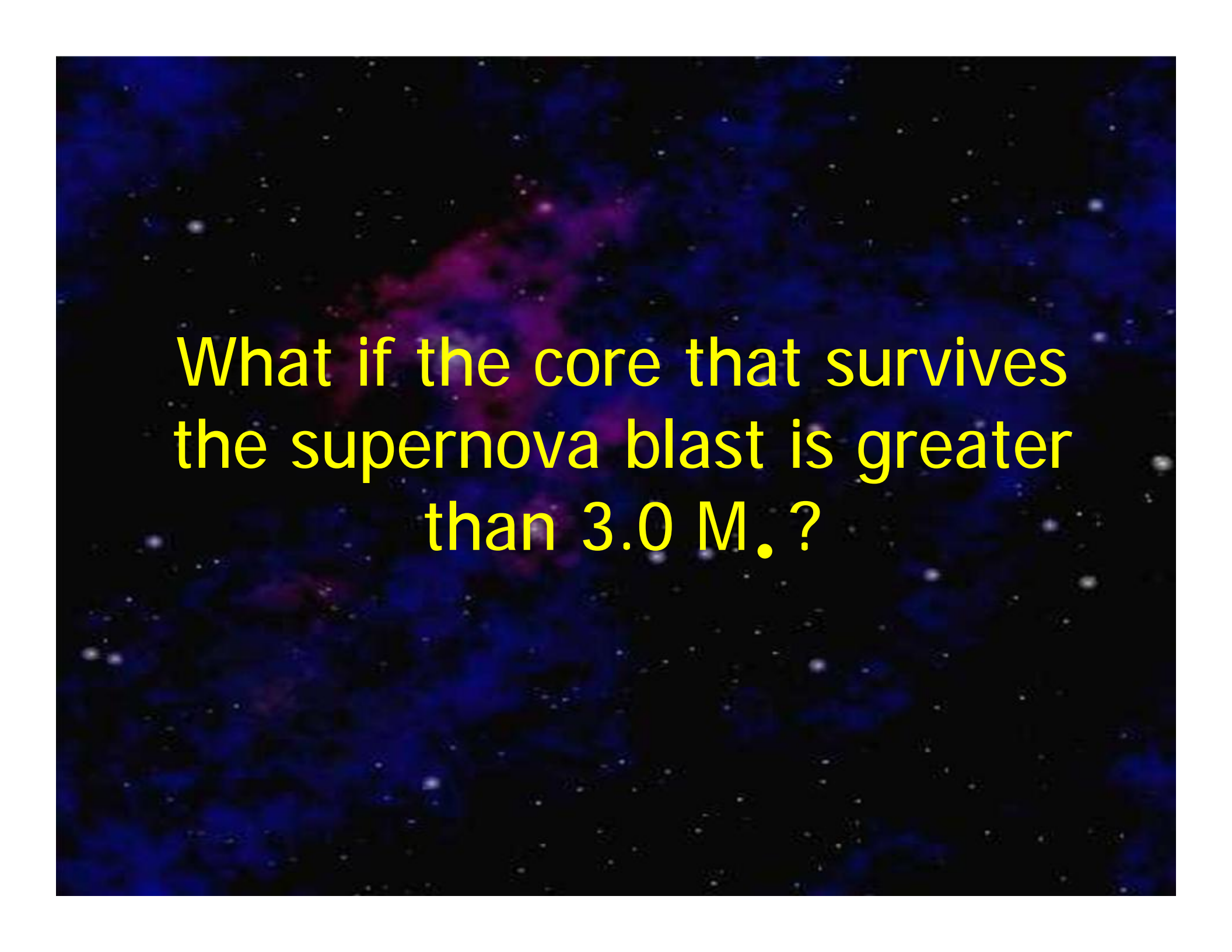


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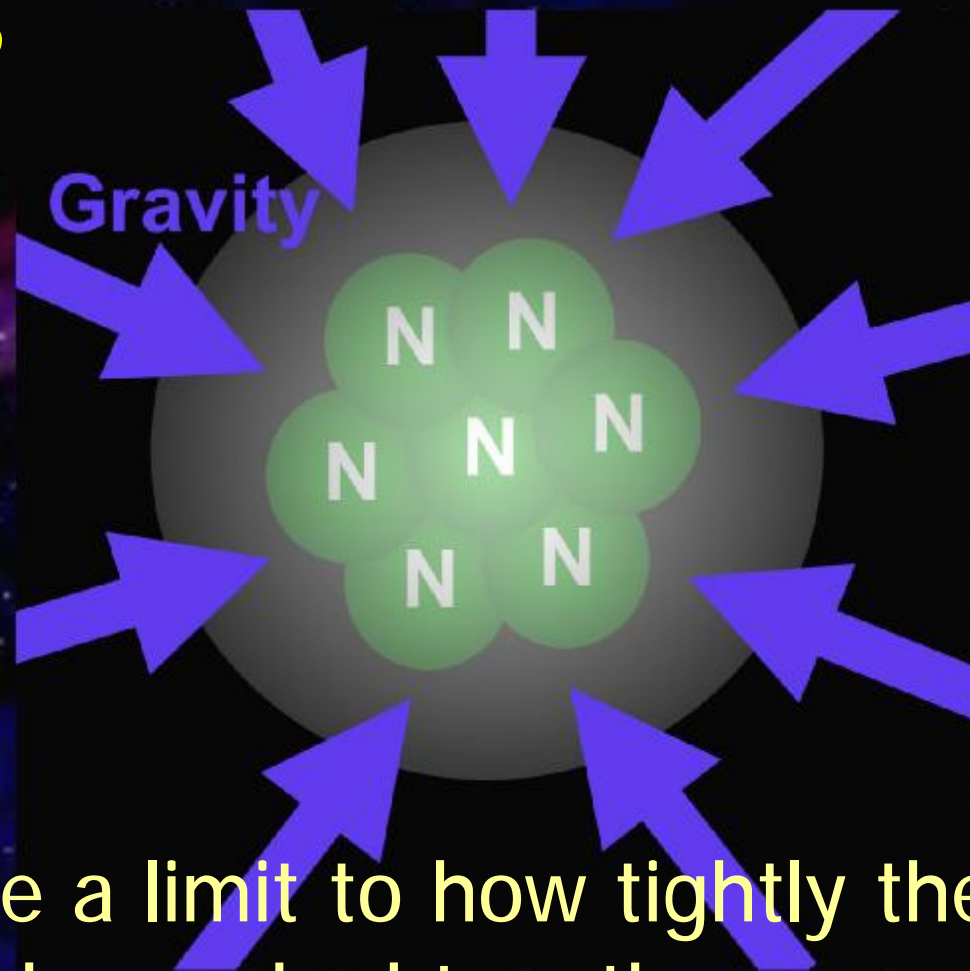






What if the core that survives
the supernova blast is greater
than 3.0 M_{\odot} ?

What keeps the neutron star from collapsing?



Neutrons have a limit to how tightly they can be packed together

NEUTRON DEGENERACY PRESSURE

Chandrasekhar Limit for neutron stars

$$M < 3.0 M_{\odot}$$



"It's black, and it looks like a hole.
I'd say it's a black hole."

What if the iron core $> 3.0M_{\odot}$.

- n Degeneracy pressure is overcome by gravity
- n The core continues to shrink producing NO HEAT.
- n No force in nature can stop the collapse

ESCAPE VELOCITY

$$V_{esc} = \sqrt{\frac{2GM}{R}}$$



G = Universal Gravitational Constant

M = Mass of the gravitating body

R = Radius of the gravitating body

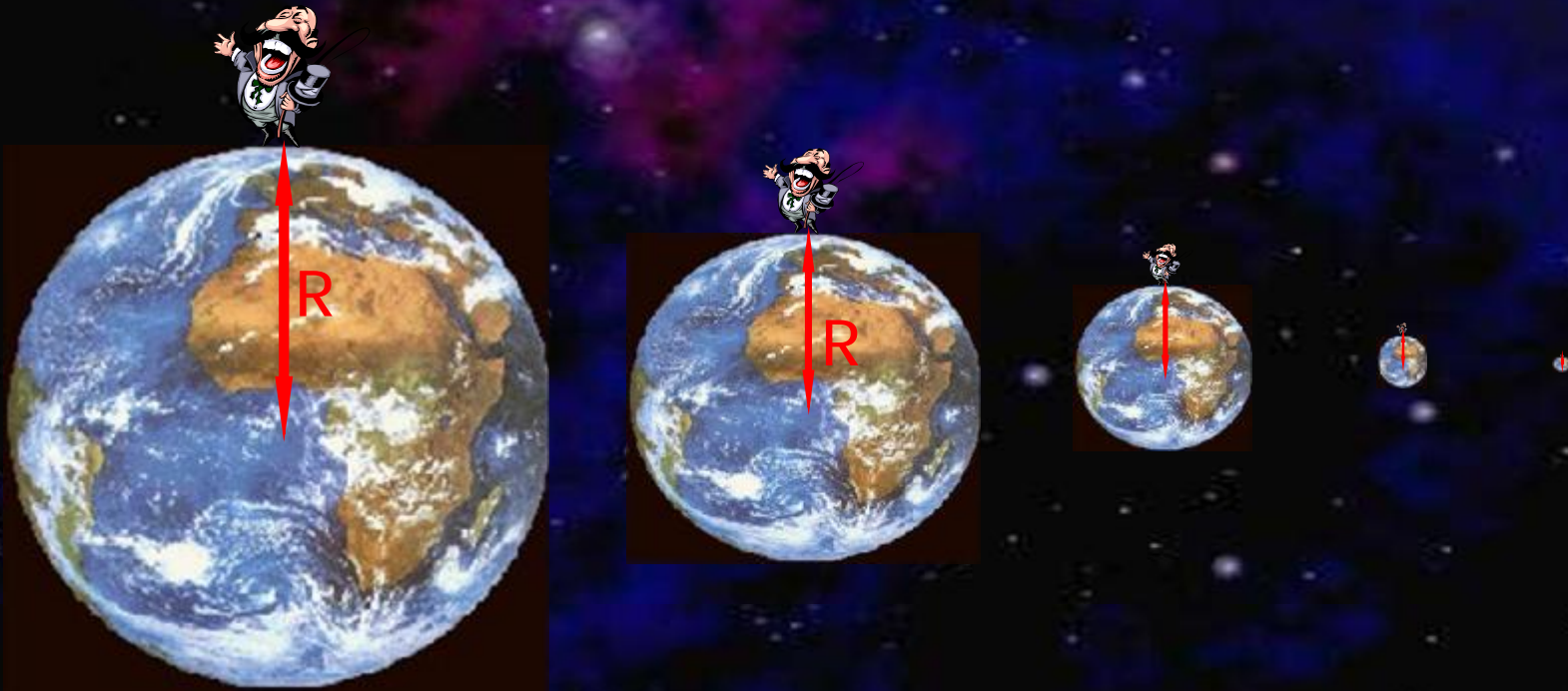
$$V_{esc} = \sqrt{\frac{2GM}{R}}$$

$$F = \frac{Gm_1m_2}{r^2}$$



$$V_{esc} = \sqrt{\frac{2GM}{R}}$$

$$F = \frac{Gm_1m_2}{r^2}$$





Soon the escape velocity is
greater than the speed of light!

Ultimately, the gravity is so strong that
NOT EVEN LIGHT CAN ESCAPE!

BLACK HOLE





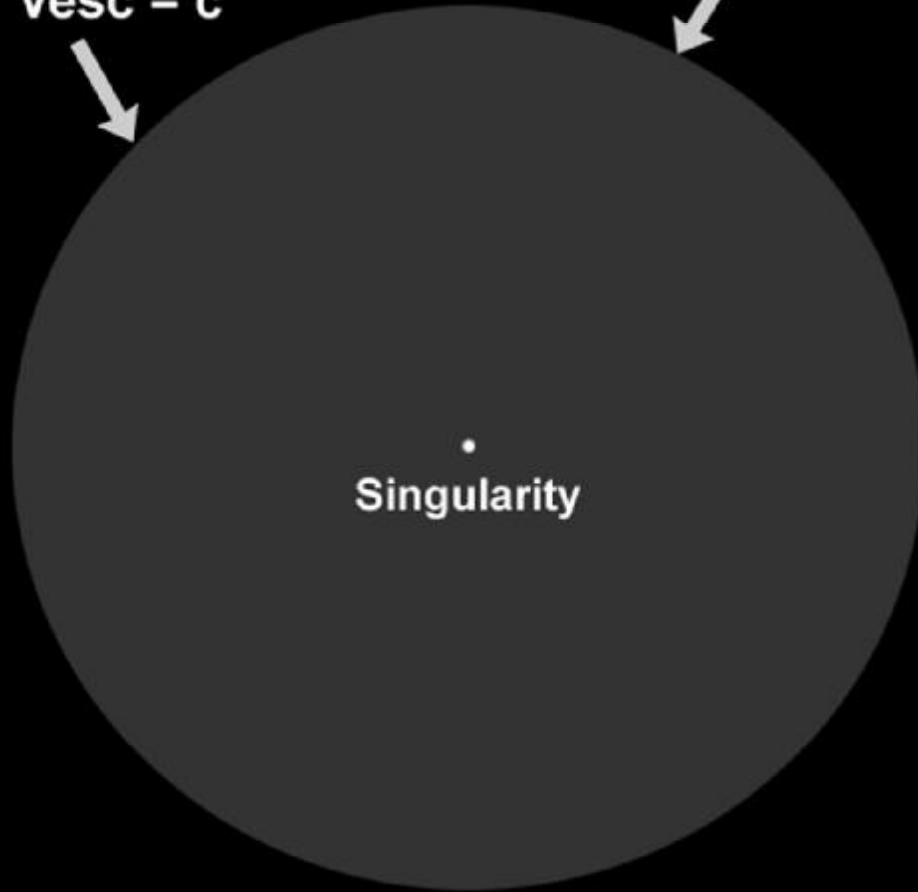
Singularity

Radius = 0

**Mass = mass of the
original core**

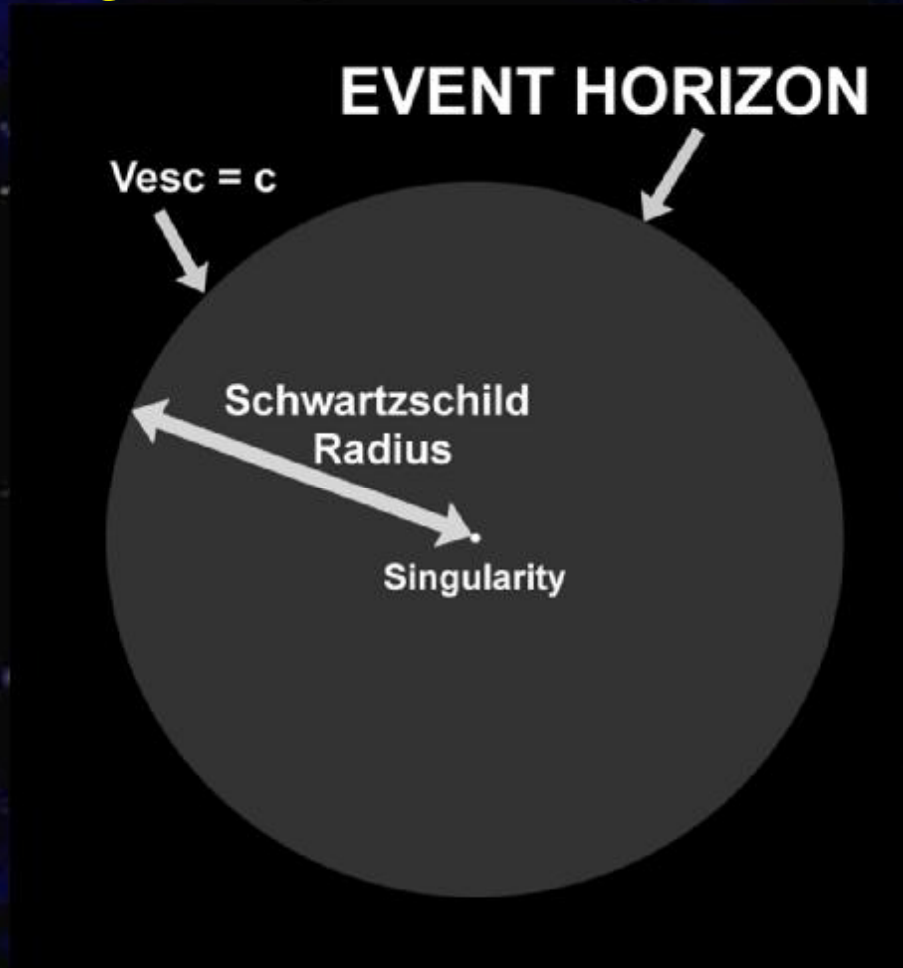
EVENT HORIZON

$V_{esc} = c$



Singularity

Physical Properties of Black Holes



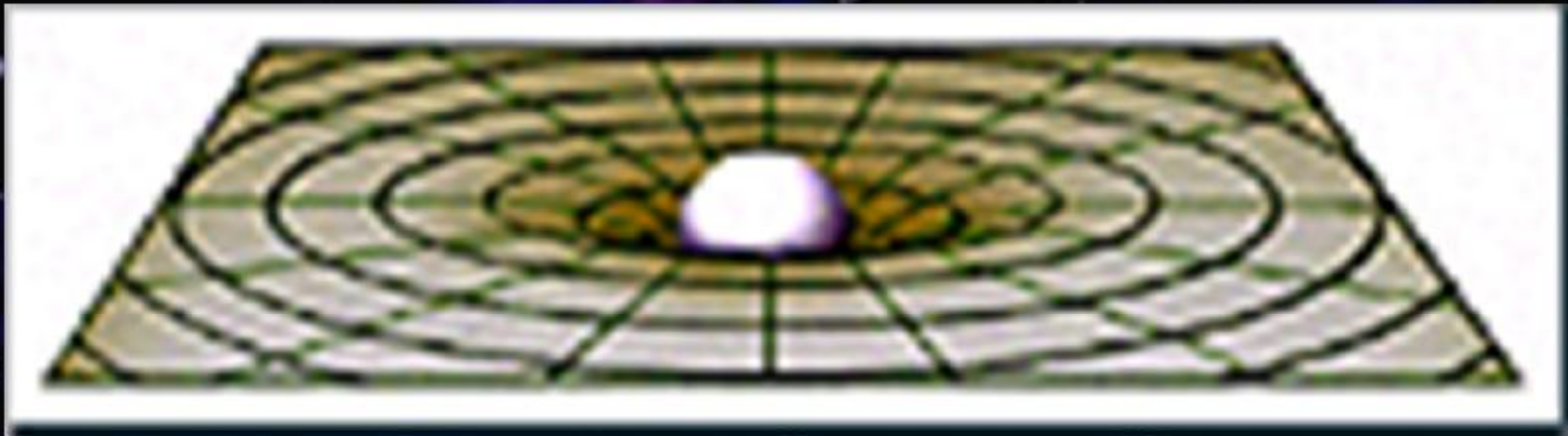
$$R = \frac{2GM}{c^2}$$

$R =$ Schwartzschild Radius

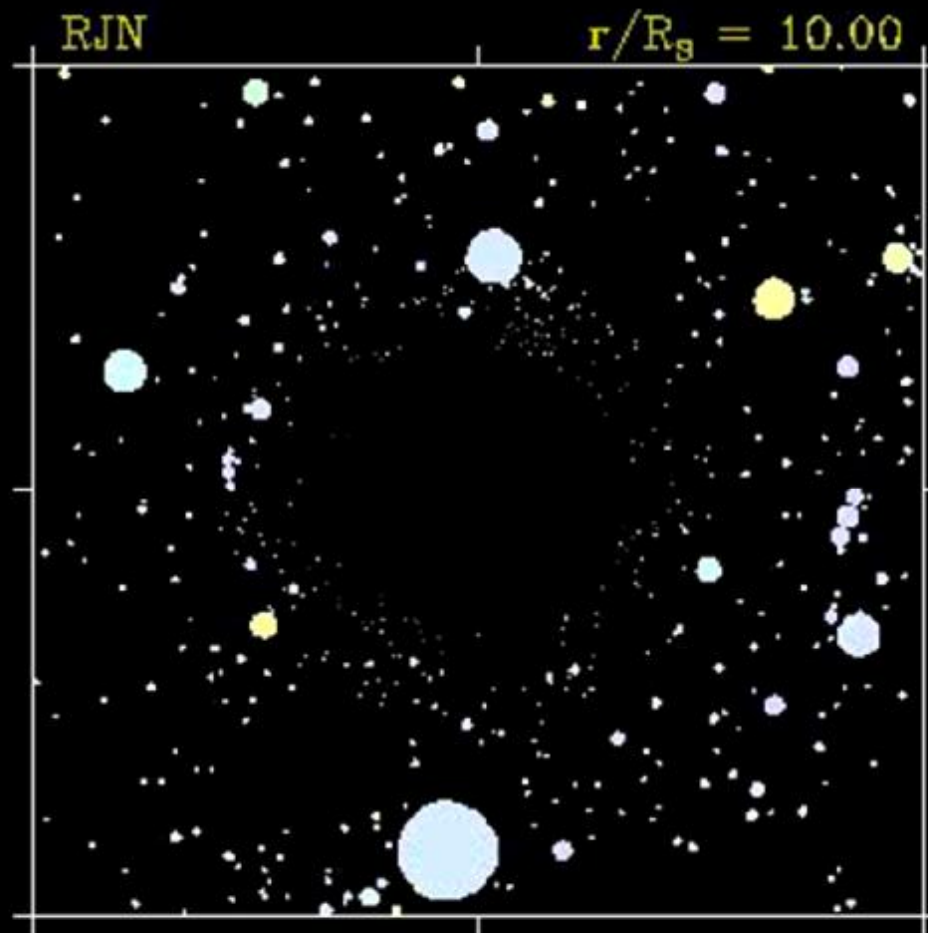
Size of event horizon depends only on MASS

Mass bends space and time

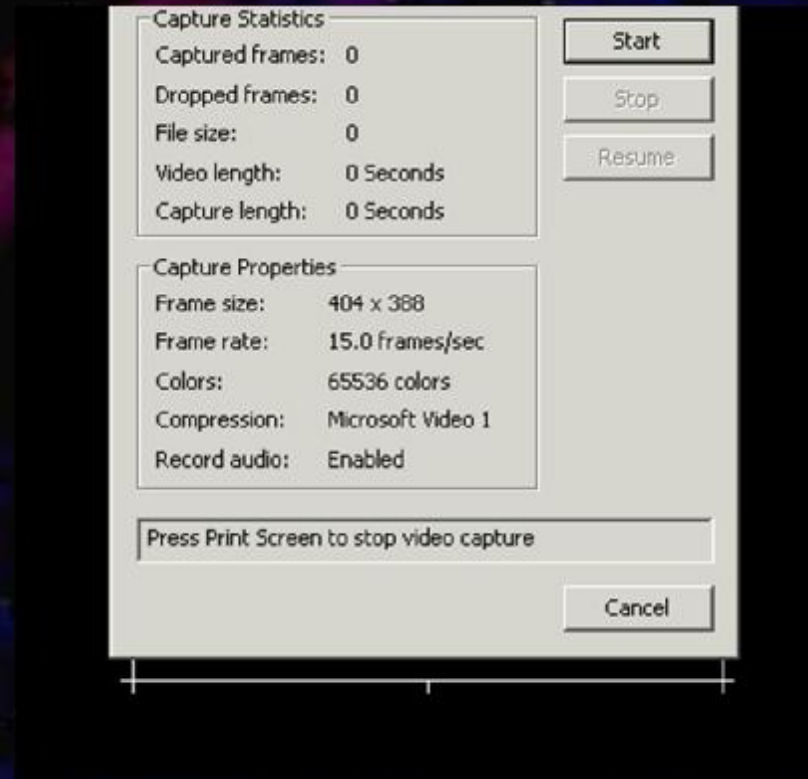
- very peculiar effects on time
- space becomes warped



In orbit around a Black Hole



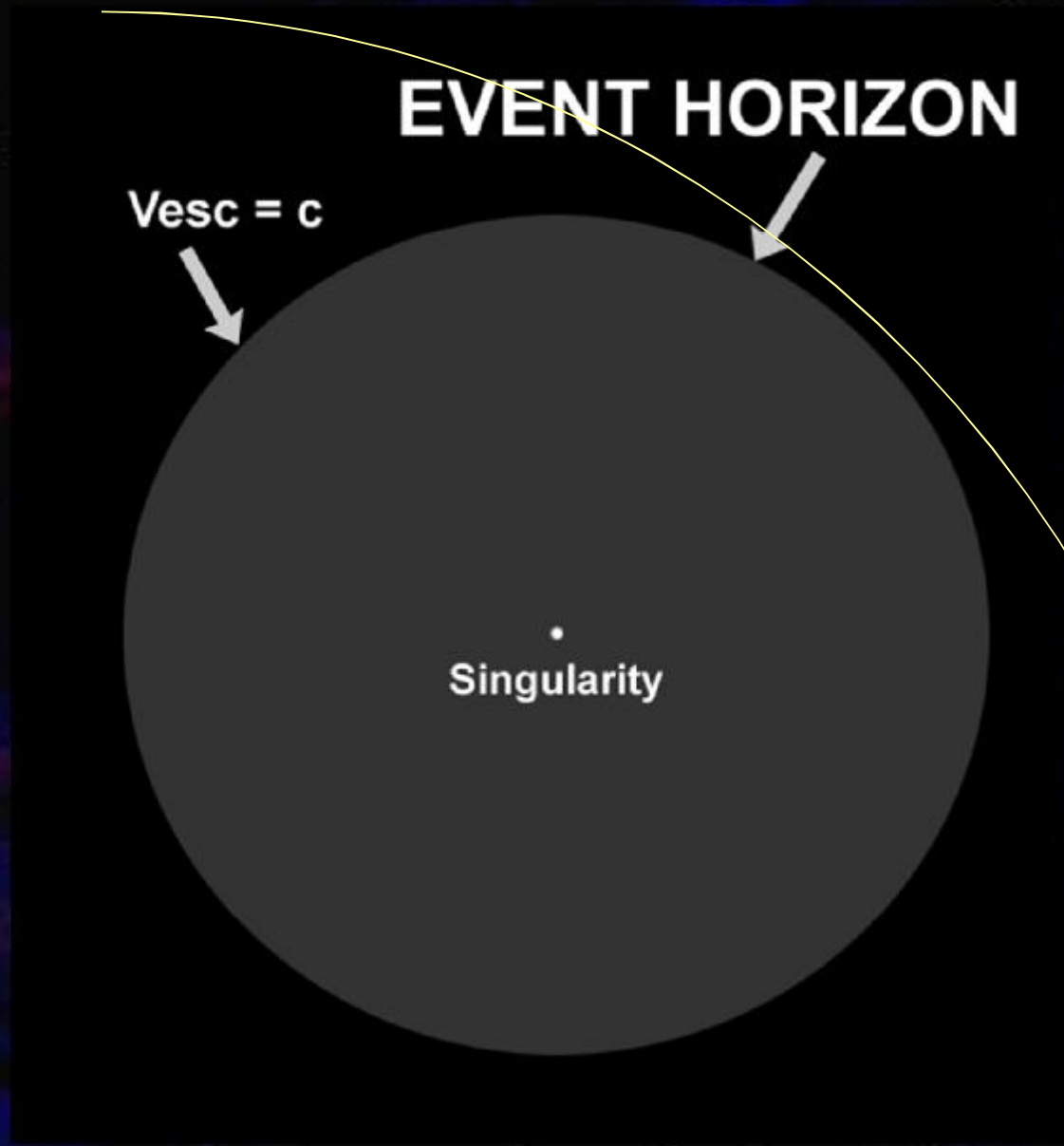
In orbit around a Black Hole at Event Horizon



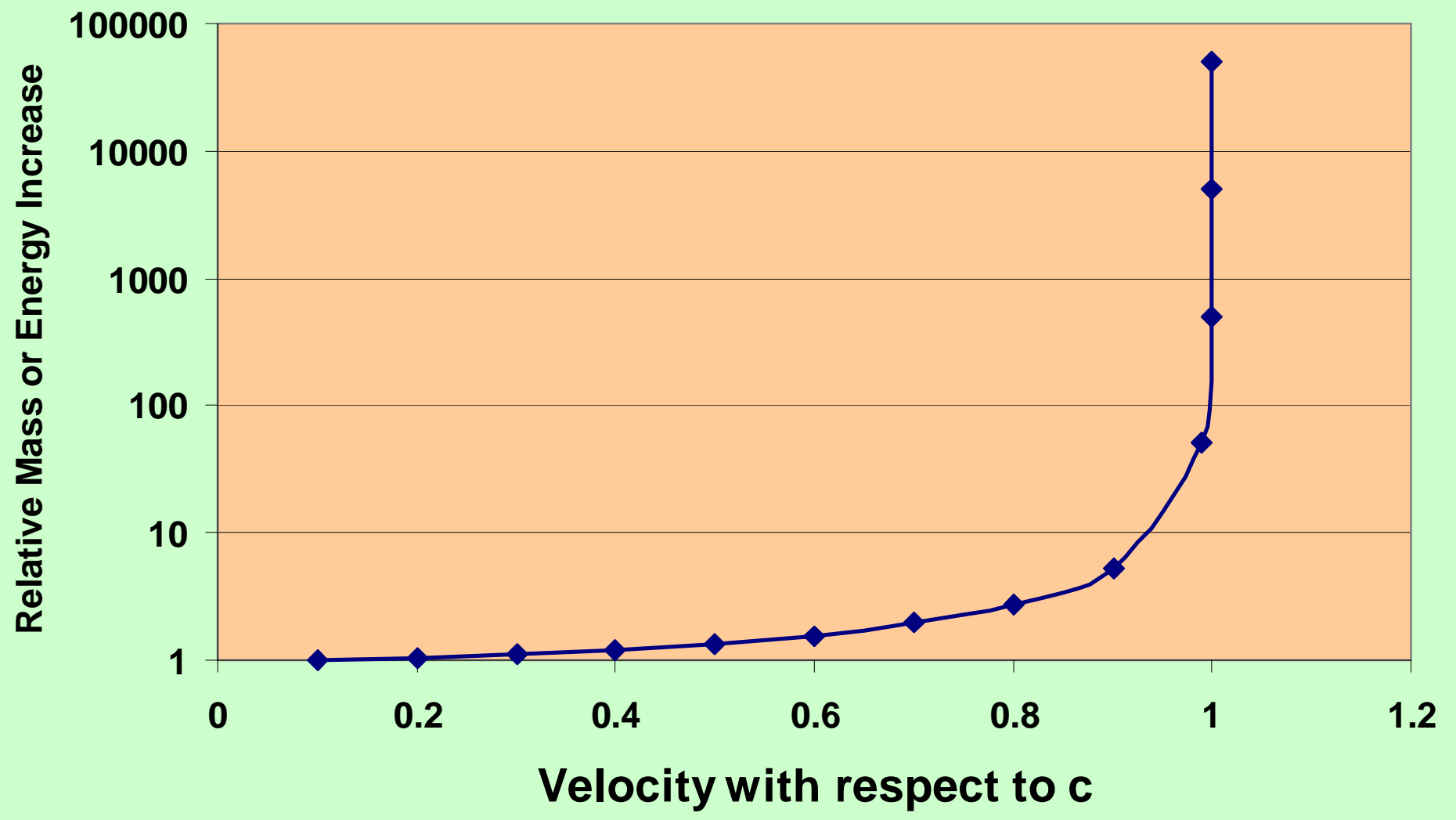
EVENT HORIZON

$V_{esc} = c$

Singularity



Effect of increasing velocity on Mass and Energy Requirements



As velocity increases
so does kinetic energy.
If $v=c$ then K goes to
infinity.

$$K = \frac{mc^2}{\sqrt{1 - (v/c)^2}} - mc^2$$

As velocity increases
so does m_{moving} . If $v=c$
then m_{moving} goes to
infinity.

$$m_{\text{moving}} = \frac{m_{\text{rest}}}{\sqrt{1 - (v/c)^2}}$$

If object has mass of 1 kg, what is its mass at 80% c

$$1 - (.8/1)^2 = 0.36 \quad \text{Sqrt } 0.36 = 0.6 \quad 1\text{kg}/0.6 = 1.666\text{kg}$$

RESULTS:

Nothing with mass can
travel AT the speed of light

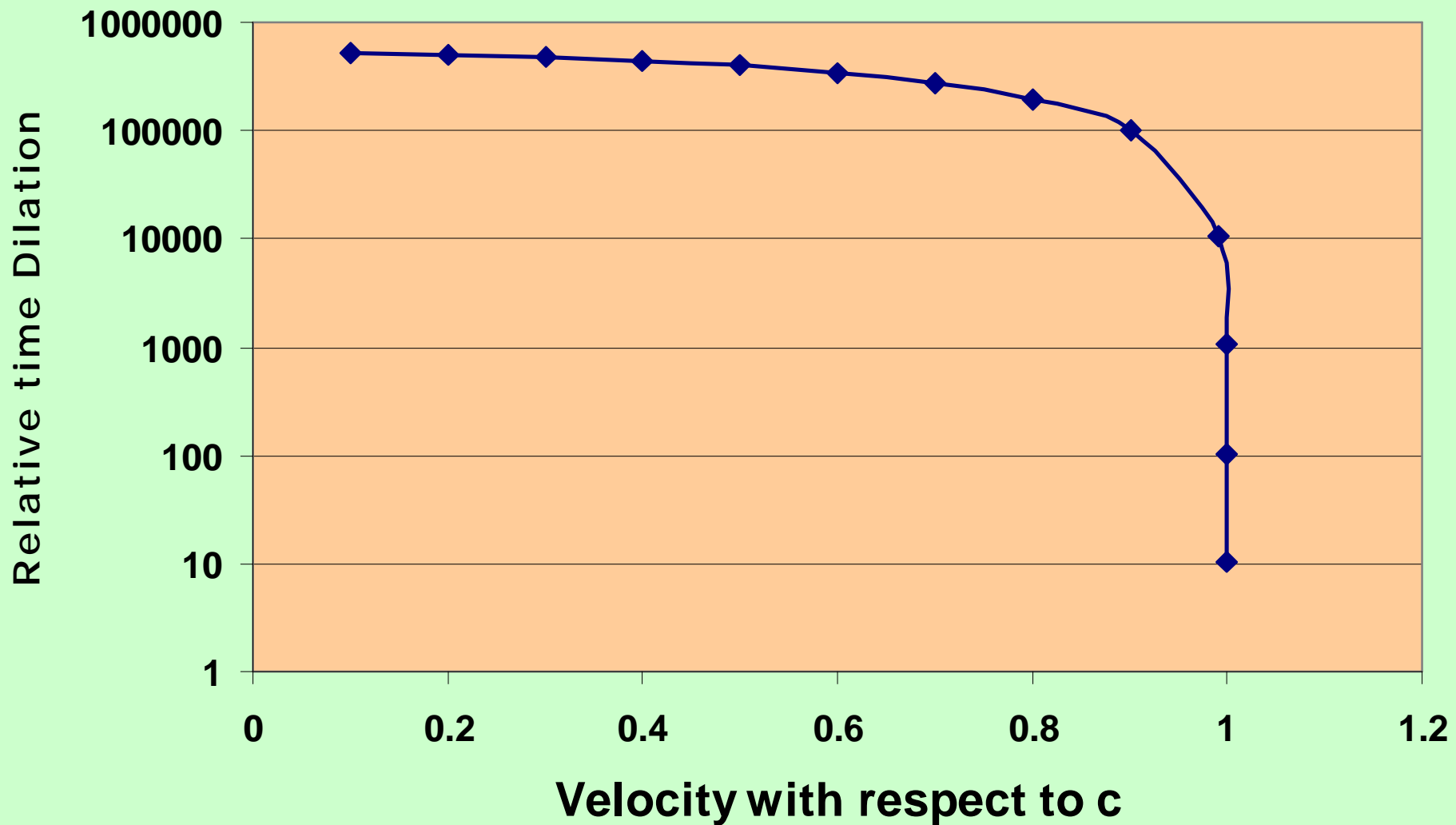
n However NEAR light speed is possible

Advantages of NEAR light speed travel: Relativistic Time Dilation

- n The measurement of the passage of time is relative to the frame of reference
- n The passage of time for someone moving at high speeds appears slower as seen by an observer at rest

$$\Delta t_{rest} = \frac{\Delta t_{moving}}{\sqrt{1 - (v/c)^2}}$$

Time Dilation as velocity increases, 1 year = 525600 minutes



50% light speed

At 50% c

- n 1.15 seconds on earth pass for every 1 second measured by a traveler
- n A 10 lightyear journey would take 20 earth years
- n Travelers would experience a 17 .4 year journey

75% light speed

- n 1.5 seconds on earth pass for every 1 second measured by a traveler
- n A 10 lightyear journey would take 13 earth years
- n Travelers would experience a 8.7 year journey

99% light speed

At 99% c

- n 7 seconds on earth pass for every 1 second measured by a traveler
- n A 10 lightyear journey would take 10.1 earth years
- n Travelers would experience a 1 year 5 month journey

99.99% light speed

At 99.99% c

- n 71 seconds on earth pass for every 1 second measured by a traveler
- n A 10 lightyear journey would take 10 earth years
- n Travelers would experience a 1 month 2 week journey

Finding the Invisible

The Voyage

Up Close and Personal



Welcome to your interstellar spacecraft. To begin your journey to the black hole, you must first escape Earth's gravity. Please select a speed.

75 m.p.h.

717 m.p.h. (the speed of sound)

25,000 m.p.h.

Examples

$$M = 3 M_{\odot}$$

$$R_S = 9\text{km} (5.4 \text{ mi})$$

$$M = 1 M_{\odot}$$

$$R_S = 3\text{km} (1.8 \text{ mi})$$

$$M = 1 M_{\text{earth}}$$

$$R_S \sim 1 \text{ cm}$$



**If we can't see 'em, how do
we find 'em?**

Solitary stellar mass black hole





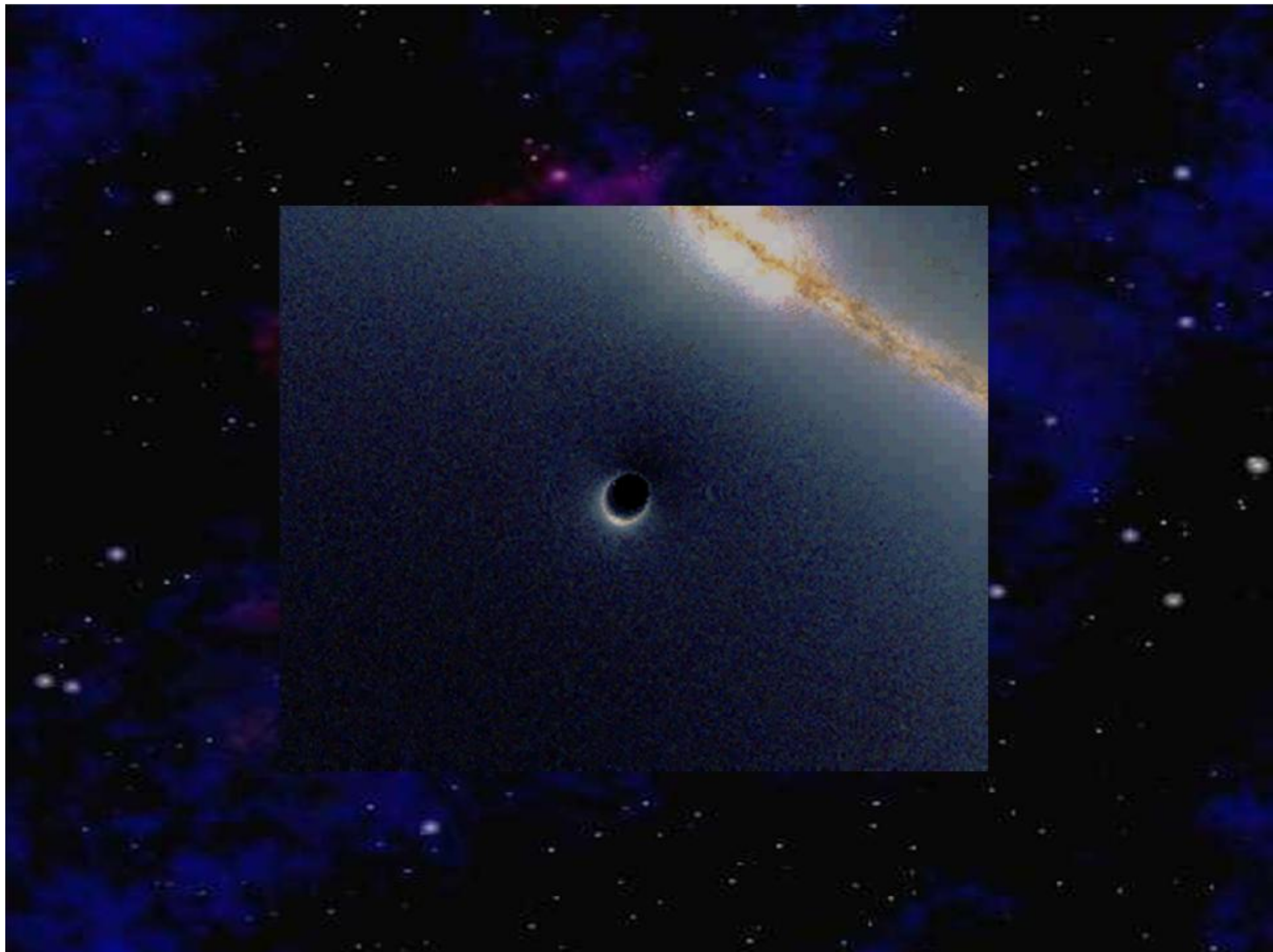
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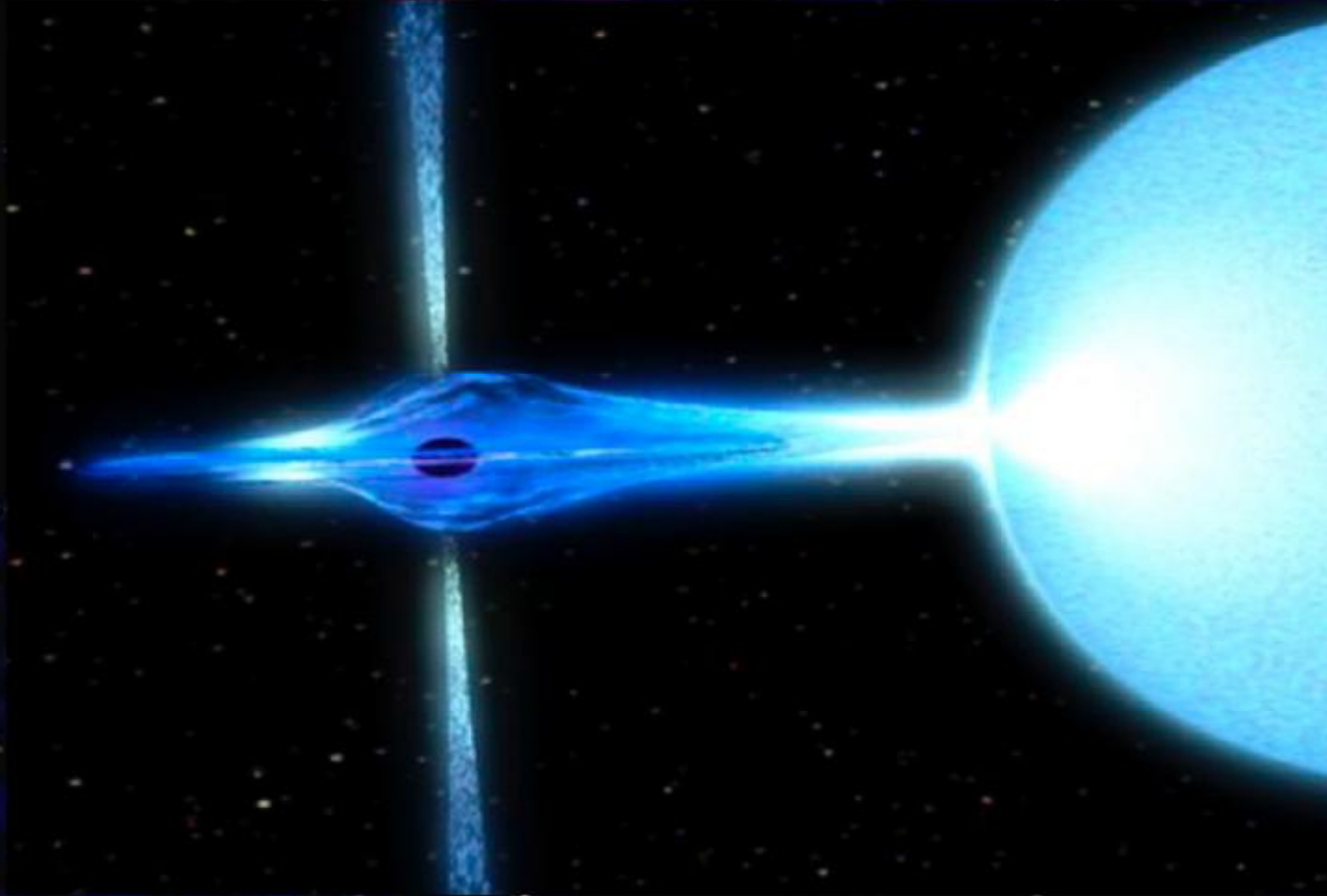
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Solitary stellar mass black hole

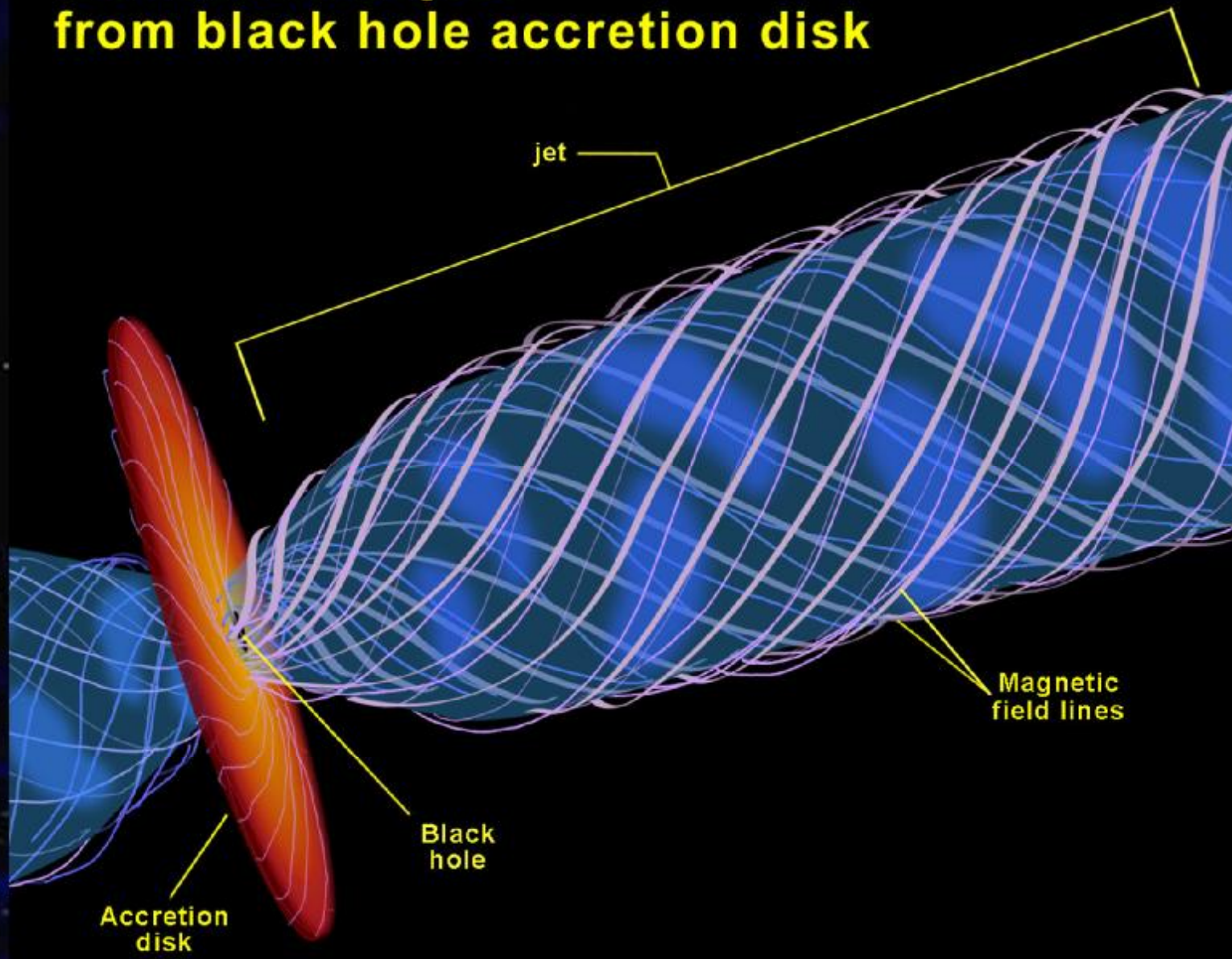




Black hole in a binary system

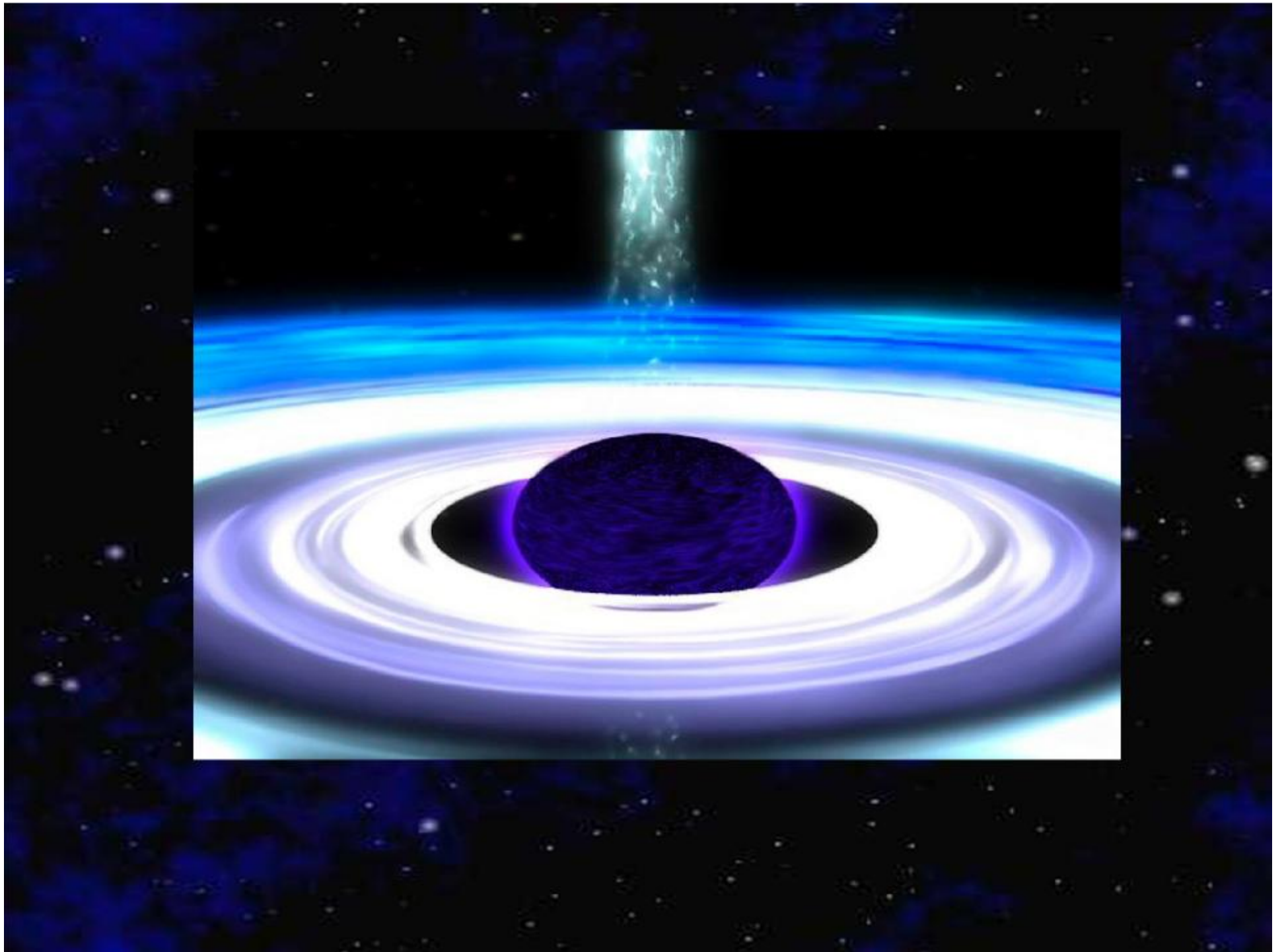


Formation of jets from black hole accretion disk



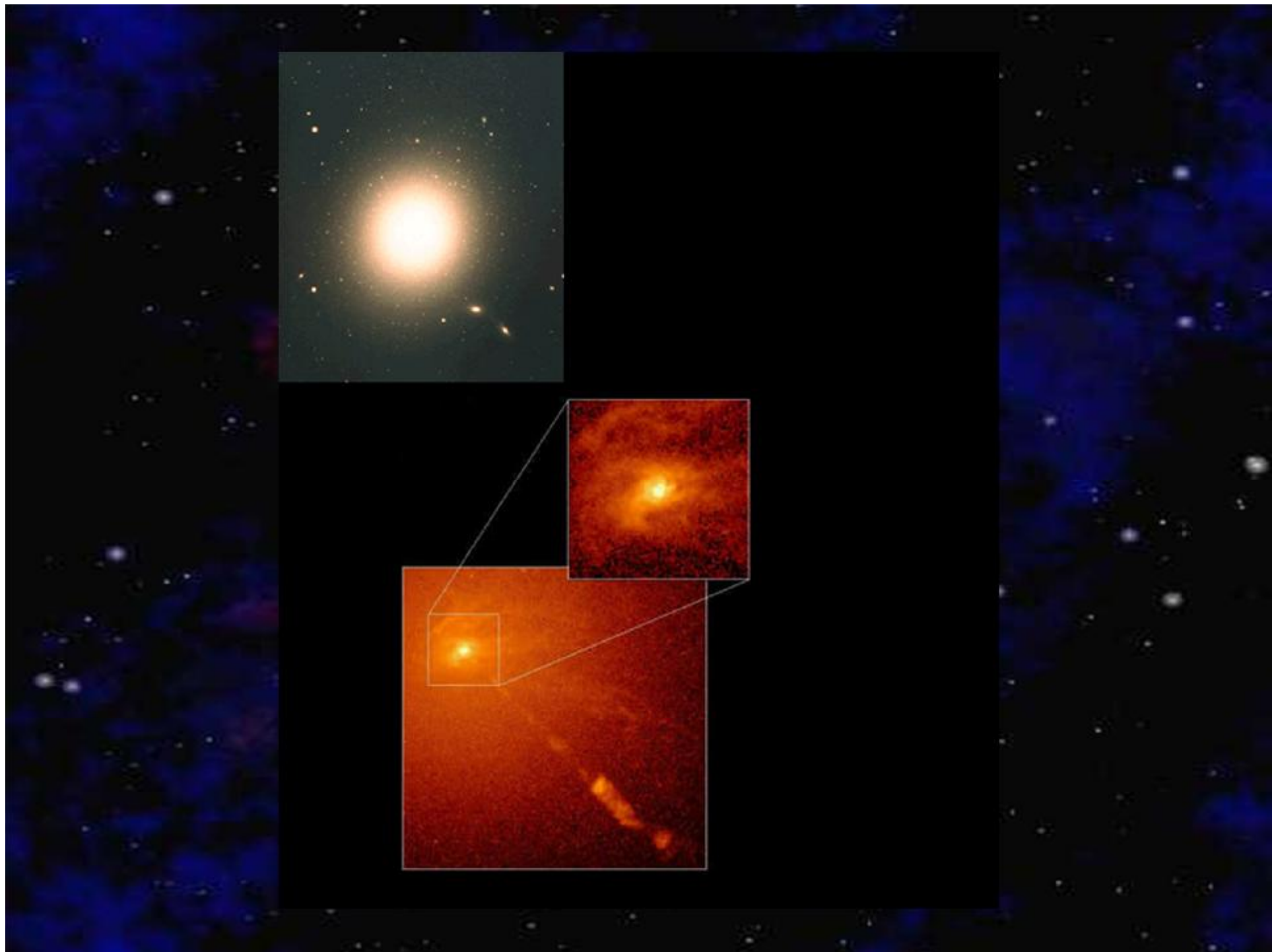


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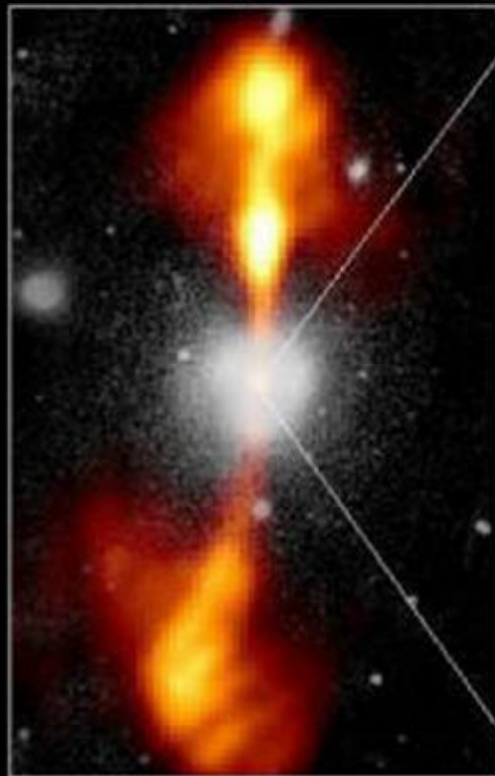


Super-massive black holes in galaxies



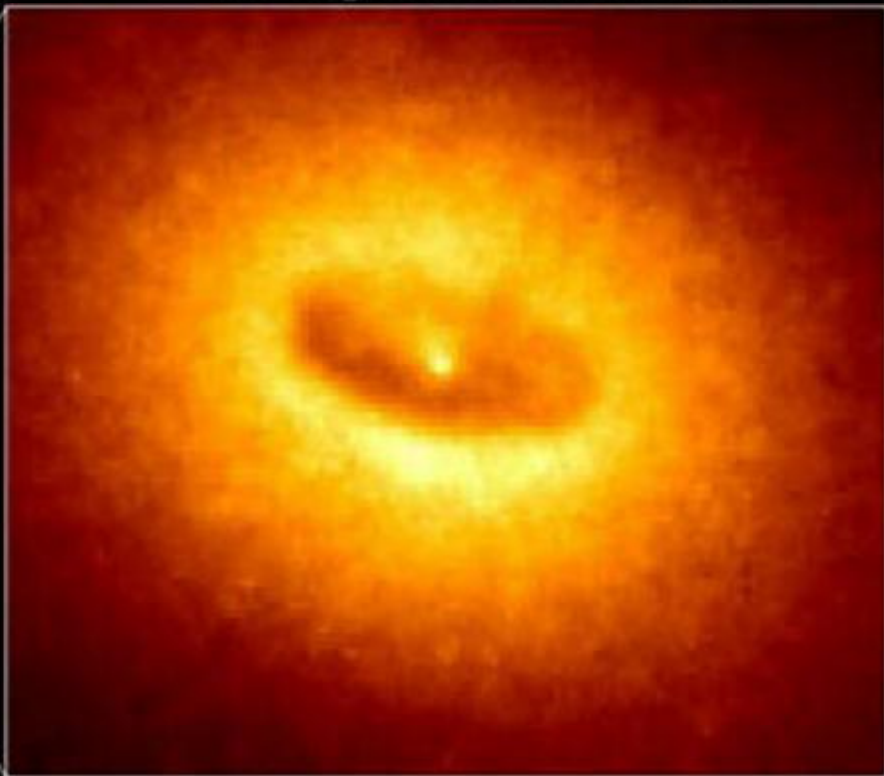


Ground-Based Optical/Radio Image



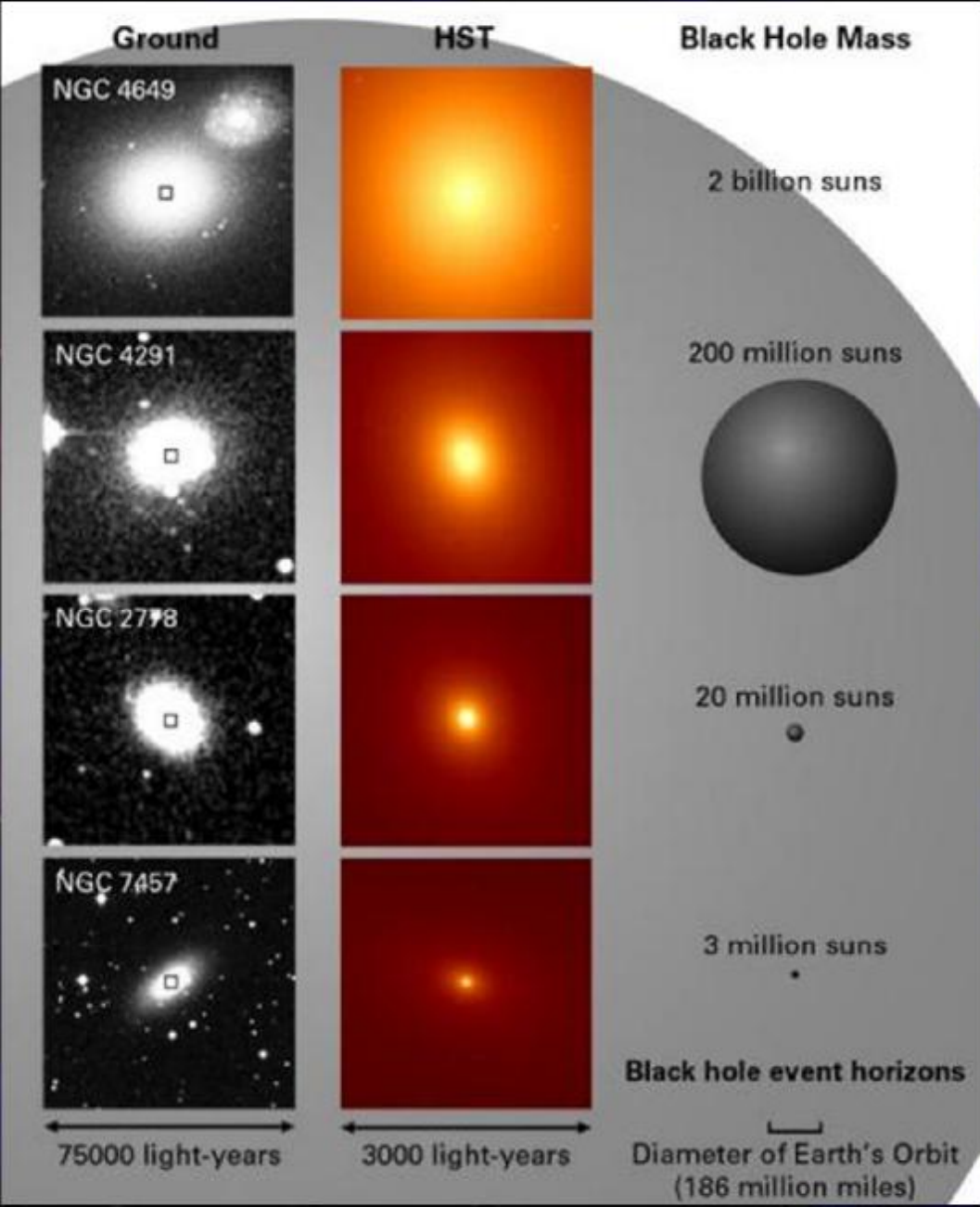
380 Arc Seconds
88,000 LIGHTYEARS

HST Image of a Gas and Dust Disk



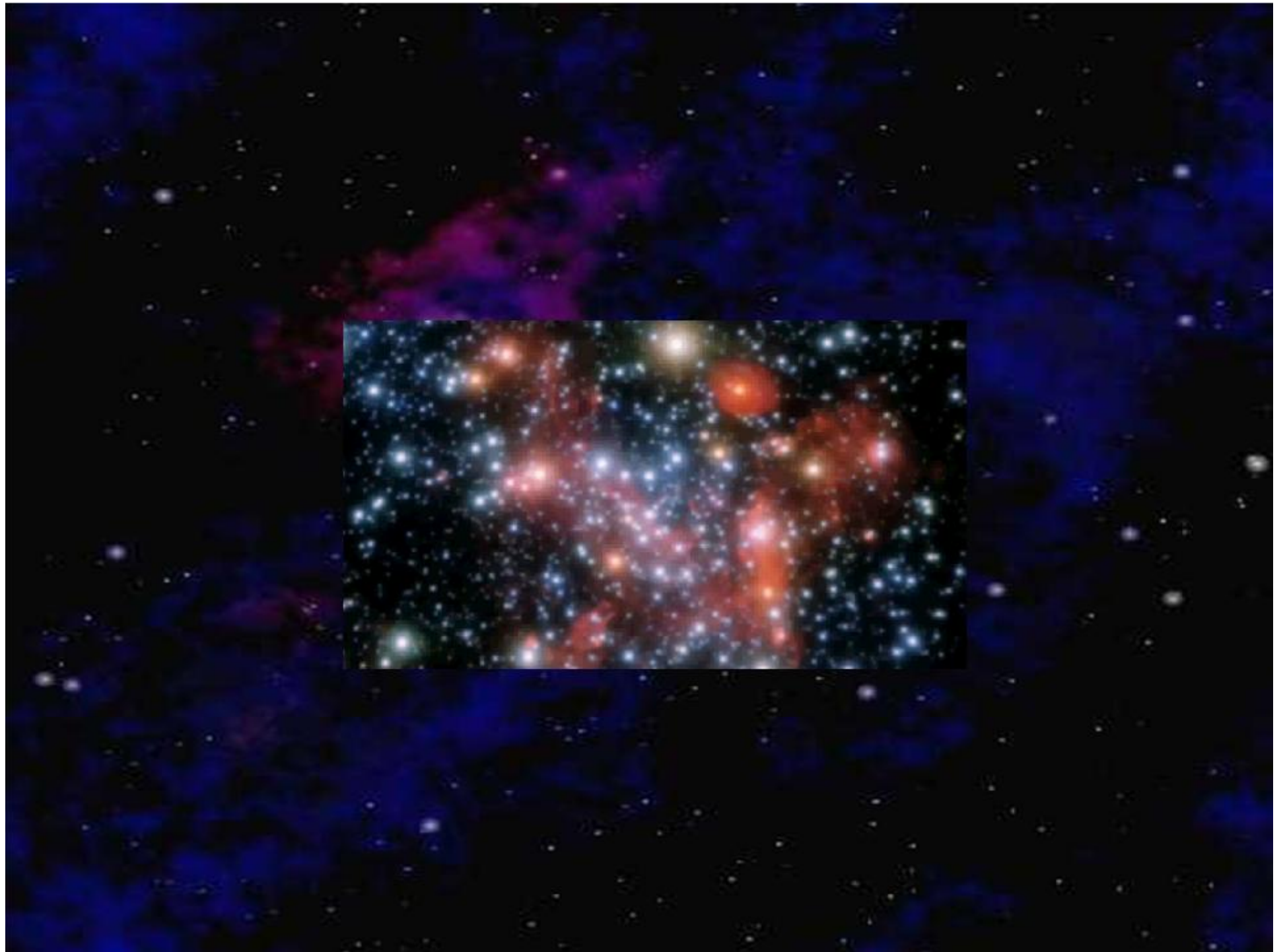
17 Arc Seconds
400 LIGHTYEARS





Probing the Galactic Center



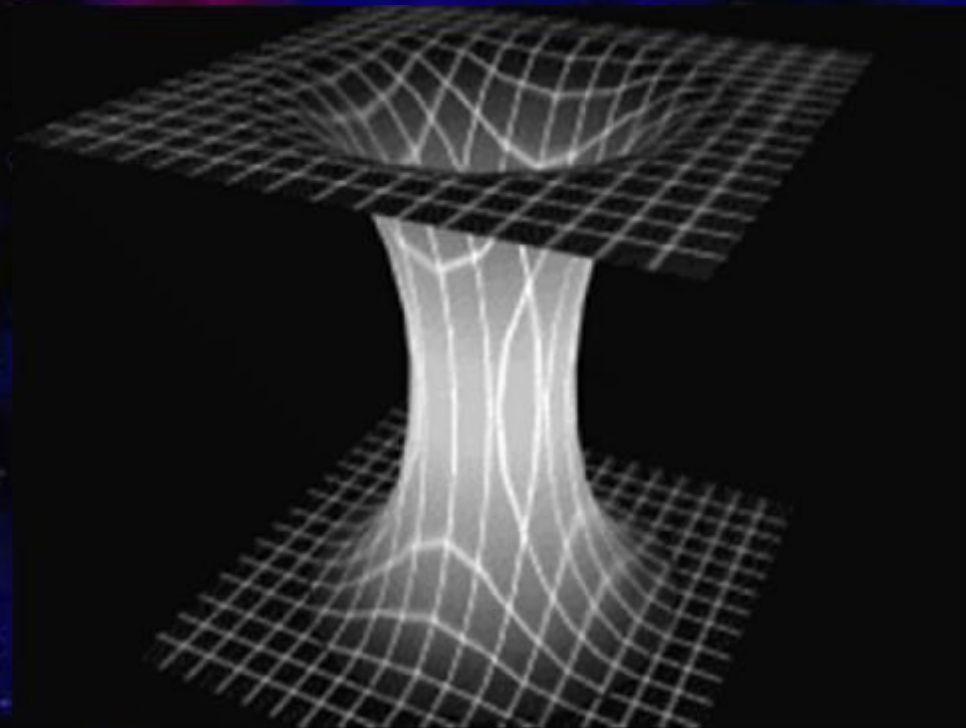


Common misconceptions about black holes:

- n Not "vacuum cleaners"
- n Not normally made from just anything

Common misconceptions about black holes:

n No "worm holes"



Common misconceptions about black holes:

n Enter at *your own risk!*

